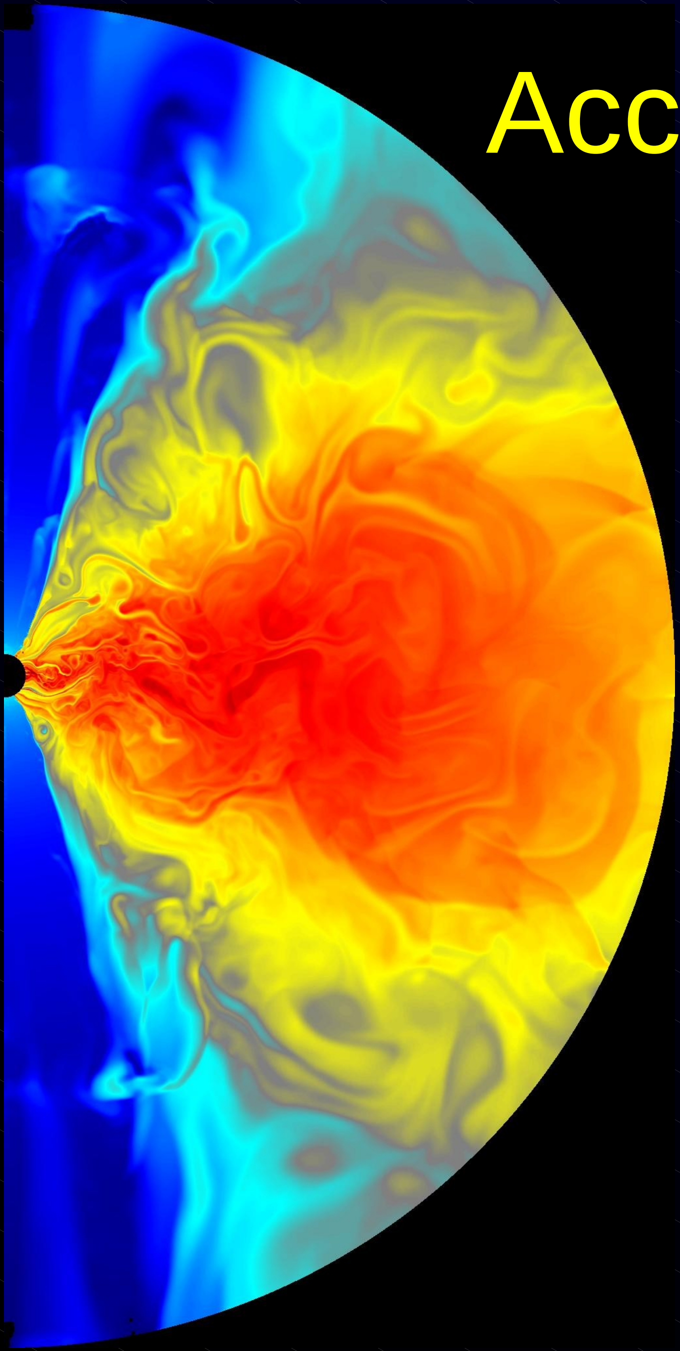
Accretion onto Black Holes



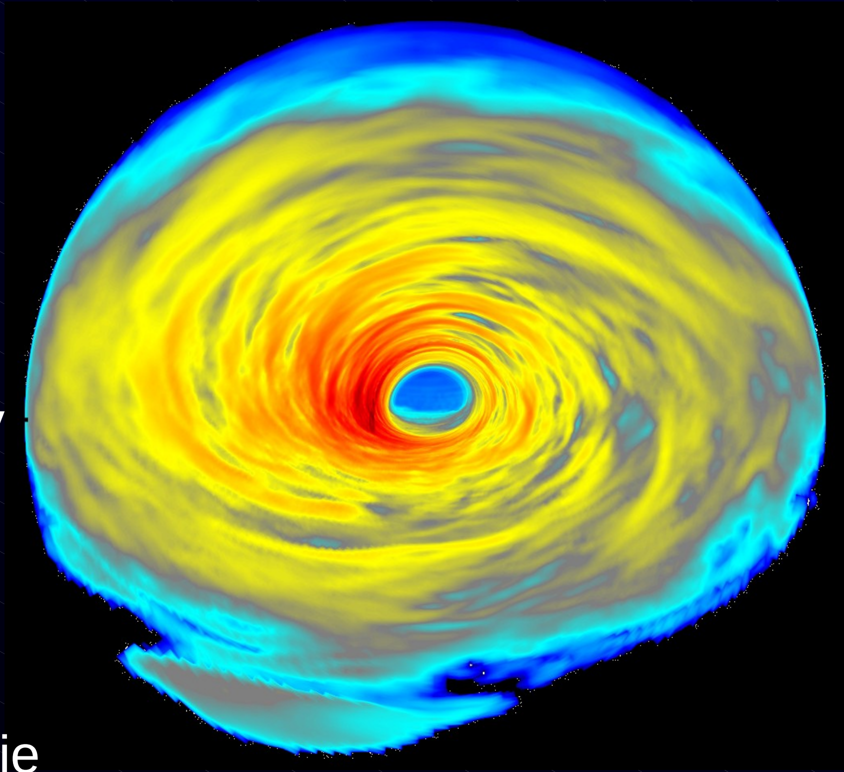
Scott C. Noble
RIT

with
Julian Krolik
JHU

John Hawley
UVa

&

Charles Gammie
UIUC



PCTS Comput. Rel. Astroph. 10/24/2009

Astrophysical Disks

Disk Type	Gravity Model
Galaxies, Stellar Disks	Newtonian
X-ray binaries, AGN, Stellar Tidal Disruptions by SMBHs	Stationary metric
Collapsars, GRBs, NS/?? Mergers, SN fall-back disks, Wet BBH Mergers	Full GR

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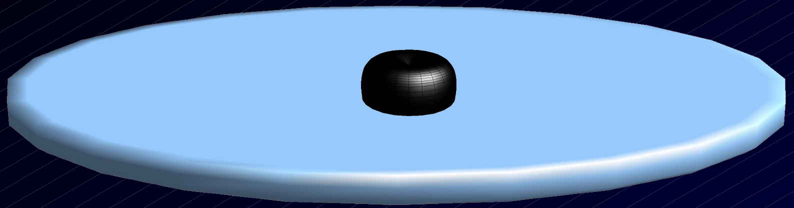
Collapsars, GRBs, NS/?? Mergers, SN fall-back disks, Wet BBH Mergers	Full GR

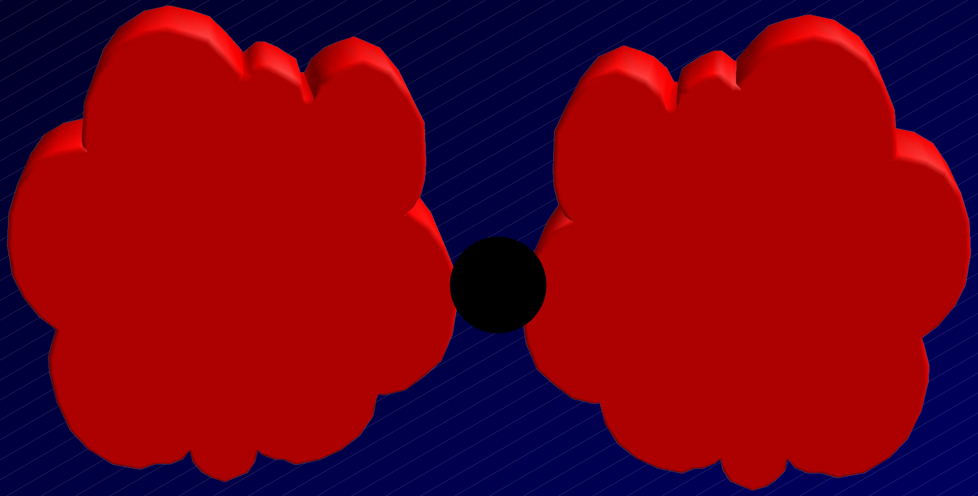
Initial Conditions Are Important!!

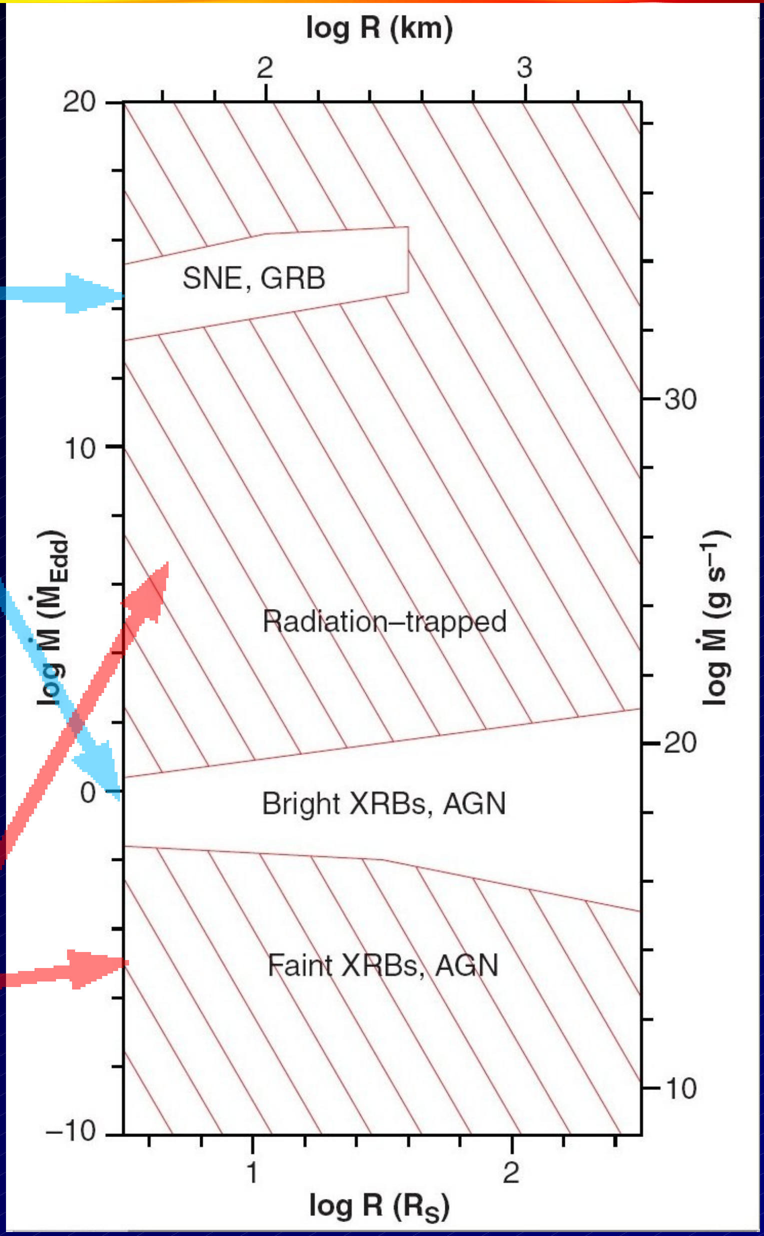
Radiative Efficiency of Disks

- Radiatively Efficient (thin disks)



- Radiatively Inefficient (thick disks)





Narayan & Quataert (2005)

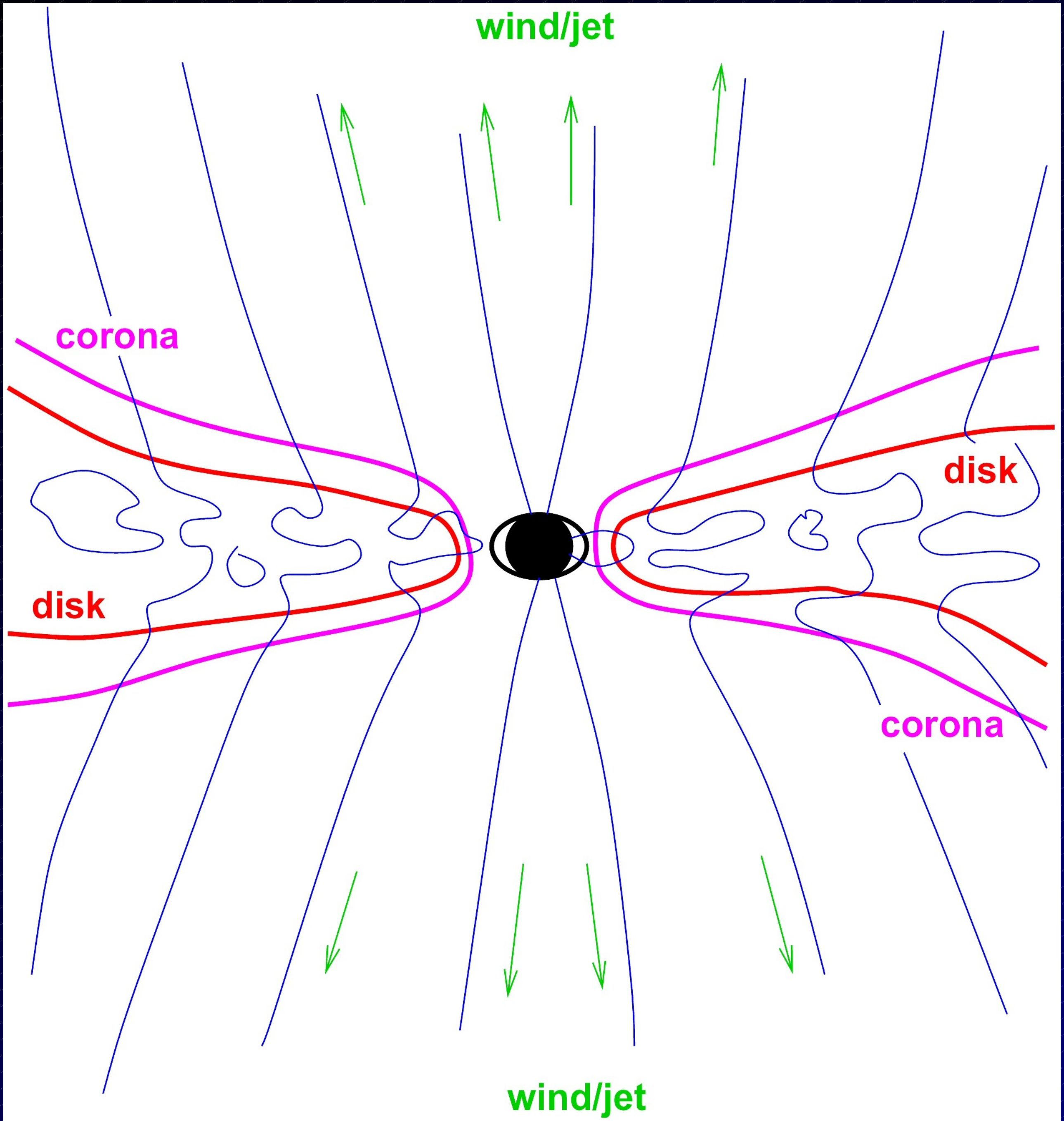


Illustration by
C. Gammie

Probing the Spacetime of BHs

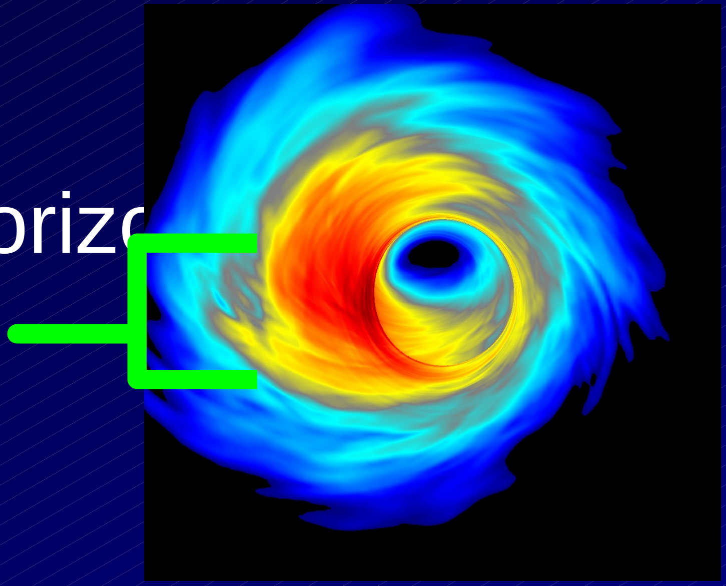
- Variability:
  - e.g. QPOs, short-time scale fluctuations

- Spectral Fitting Thermal Emission

$$L = A\, R_{in}^2\, T_{max}^4 \qquad R_{in} = R_{in}(M, a)$$

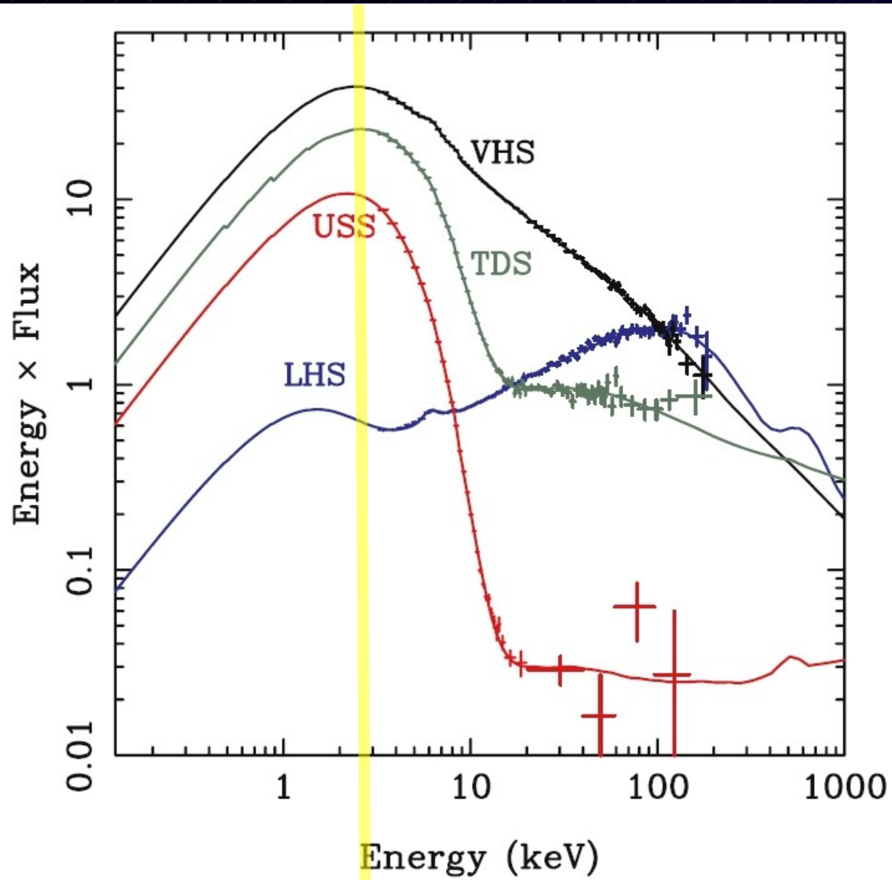
- Relativistic Iron Lines

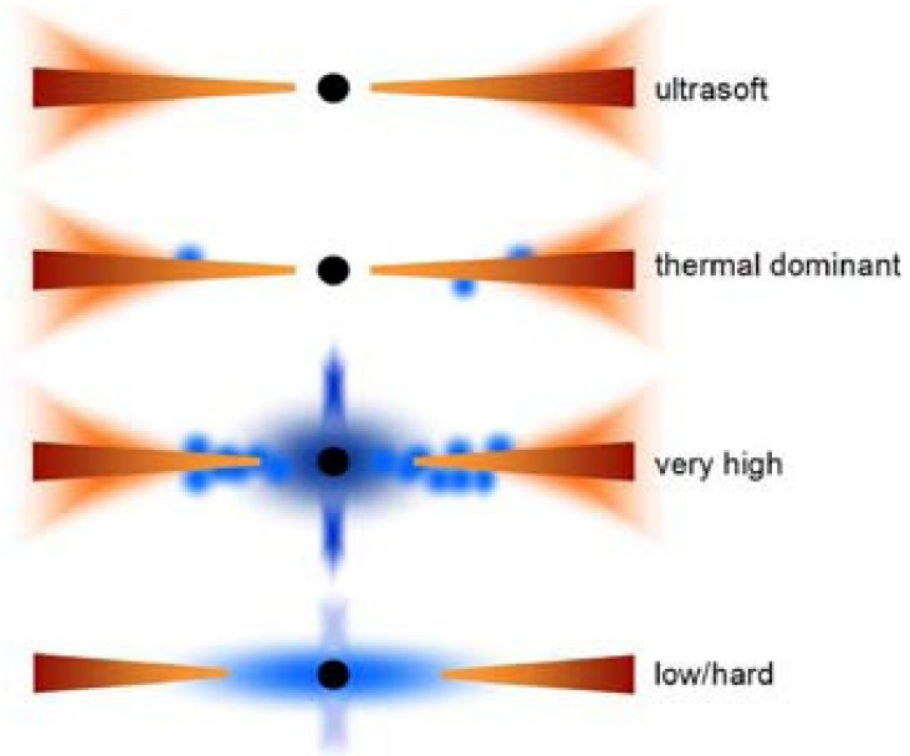
- Directly Resolving Event Horizo
  (e.g., Sgr A*)
  - Silhouette size = D(M,a)



(See Doeleman et al. (2008) for sub-mm VLBI)

Accretion States of XRBs





T_{max}

Done, Gierlinski & Kubota (2007)

$$L = A R_{in}^2 T_{max}^4$$

$$R_{in} = R_{in}(M, a) \sim R_{isco}$$

Spectral Fits for BH Spin

TABLE 1

BLACK HOLE SPIN ESTIMATES USING THE MEAN OBSERVED VALUES OF M, D, AND i

Candidate	Observation Date	Satellite	Detector	a_* (D05)	a_* (ST95)
GRO J1655−40	1995 Aug 15	*ASCA*	GIS2	~0.85	~0.8
			GIS3	~0.80	~0.75
	1997 Feb 25–28	*ASCA*	GIS2	~0.75[a]	~0.70
			GIS3	~0.75[a]	~0.7
	1997 Feb 26	*RXTE*	PCA	~0.75[a]	~0.65
	1997 (several)	*RXTE*	PCA	0.65–0.75[a]	0.55–0.65
4U 1543−47	2002 (several)	*RXTE*	PCA	0.75–0.85[a]	0.55–0.65

[a] Values adopted in this Letter.

Shafee et al. (2006)

Object	Power Law Mean	Power Law Standard Deviation
GRS 1915+105[a]	0.998	0.001
GRS 1915+105[b]	0.998	0.001

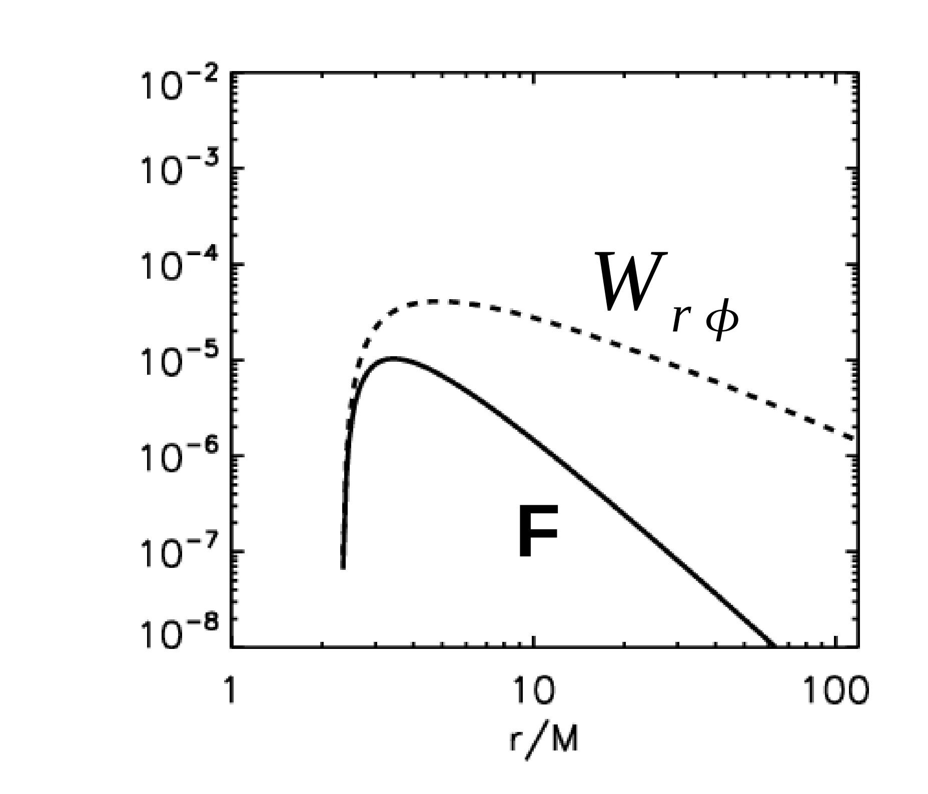
McClintock et al. (2006)

Steady-State Models: Novikov & Thorne (1973)

Assumptions:

1) Stationary gravity
2) Equatorial Keplerian Flow
   - Thin, cold disks
3) Time-independent
4) Work done by stress locally dissipated into heat
5) Conservation of M, E, L
6) Zero Stress at ISCO
   - Eliminated d.o.f.
   - Condition thought to be suspect from very start

   (Thorne 1974, Page & Thorne 1974)



$$\eta = 1 - \dot{E}/\dot{M}$$
$$= 1 - \epsilon_{ISCO}$$

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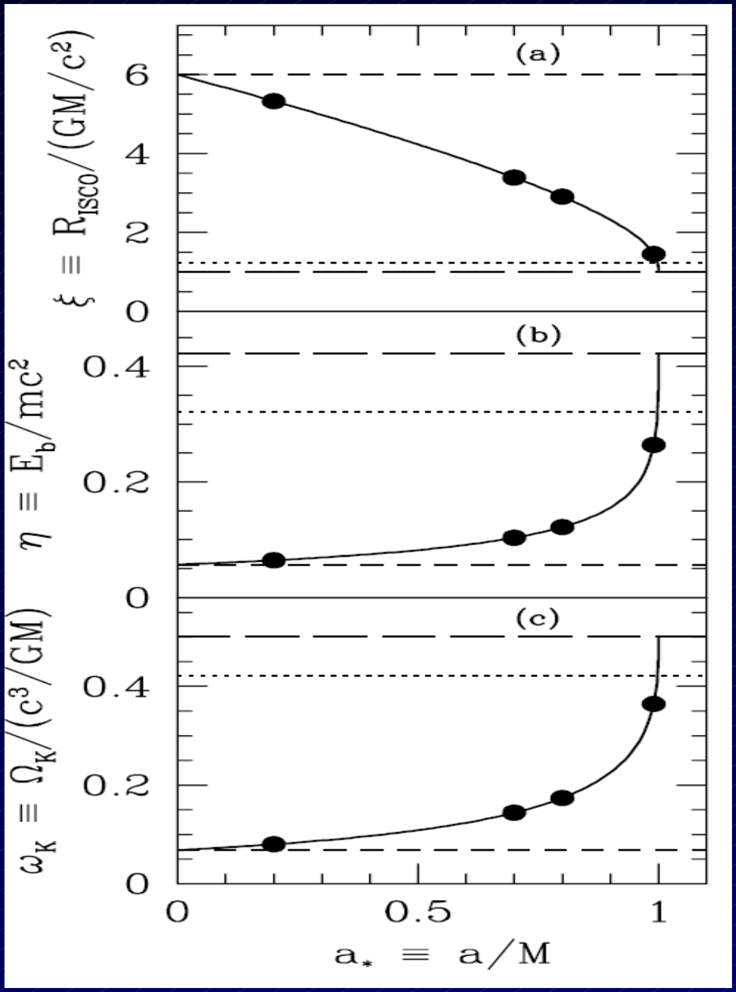
   (Thorne 1974, Page & Thorne 1974)

Magnetic Fields → Need dynamical evolution!!!



$$\eta = 1 - \dot{E}/\dot{M}$$
$$= 1 - \epsilon_{ISCO}$$

Steady-State Models: α Disks

- Shakura & Sunyaev (1973):

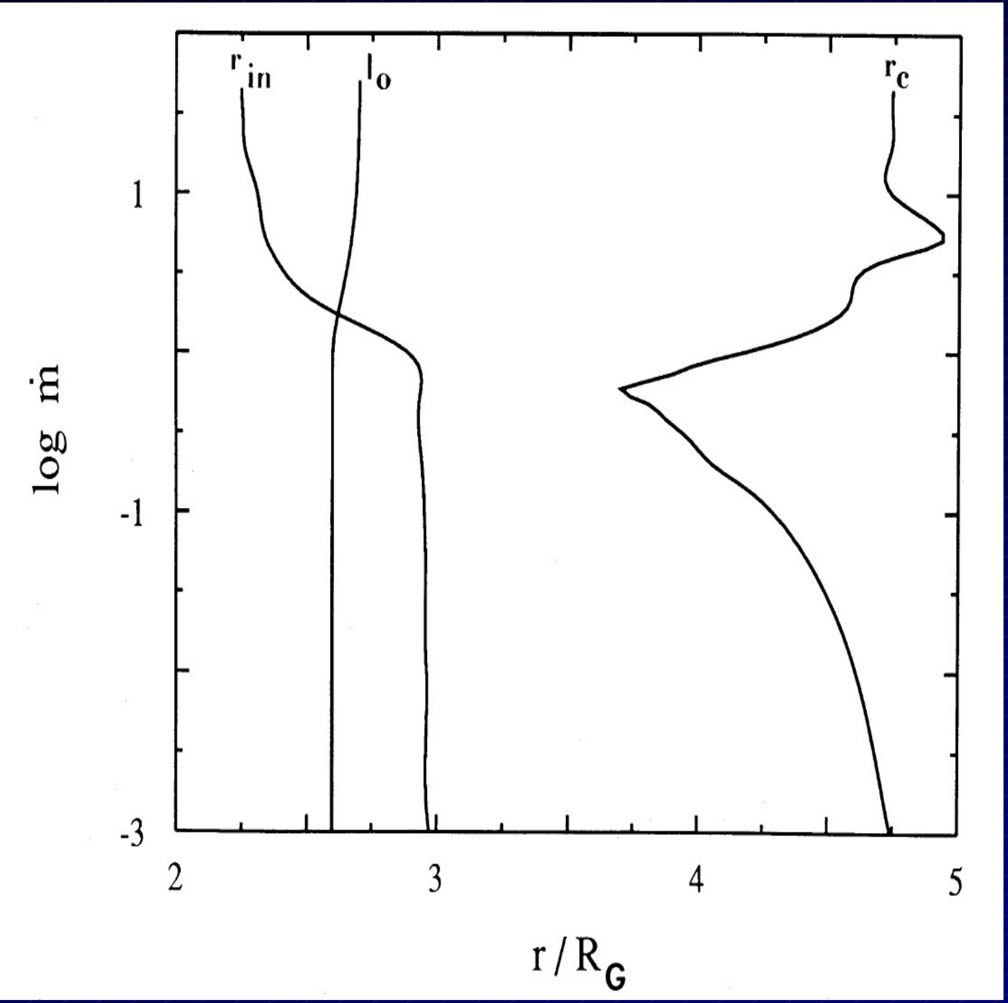
$$T^r_\phi = -\alpha P$$

$$P = \rho c_s^2 \qquad t^r_\phi = -\alpha c_s^2$$

- No stress at sonic point:
  $\rightarrow$ $R_{in} = R_s$
  e.g.:
  Muchotrzeb & Paczynski (1982)
  Abramowicz, et al. (1988)
  Afshordi & Paczyncski (2003)

  (Schwarzschild BHs)

- Variable α
  e.g., Shafee, Narayan, McClintock (2008)

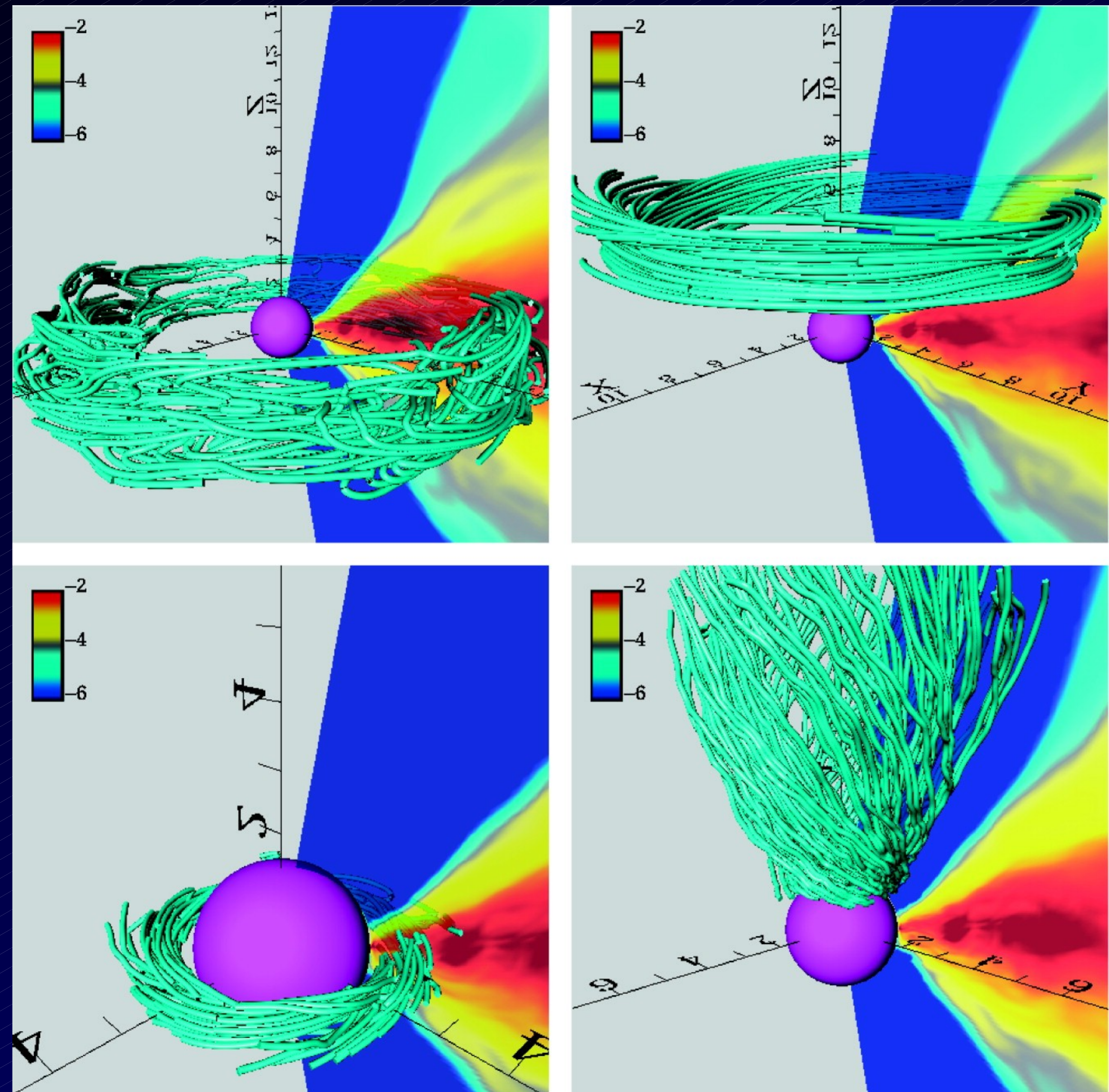


Abramowicz, et al. (1988)

$$\eta \sim 1 - \epsilon_{isco}$$

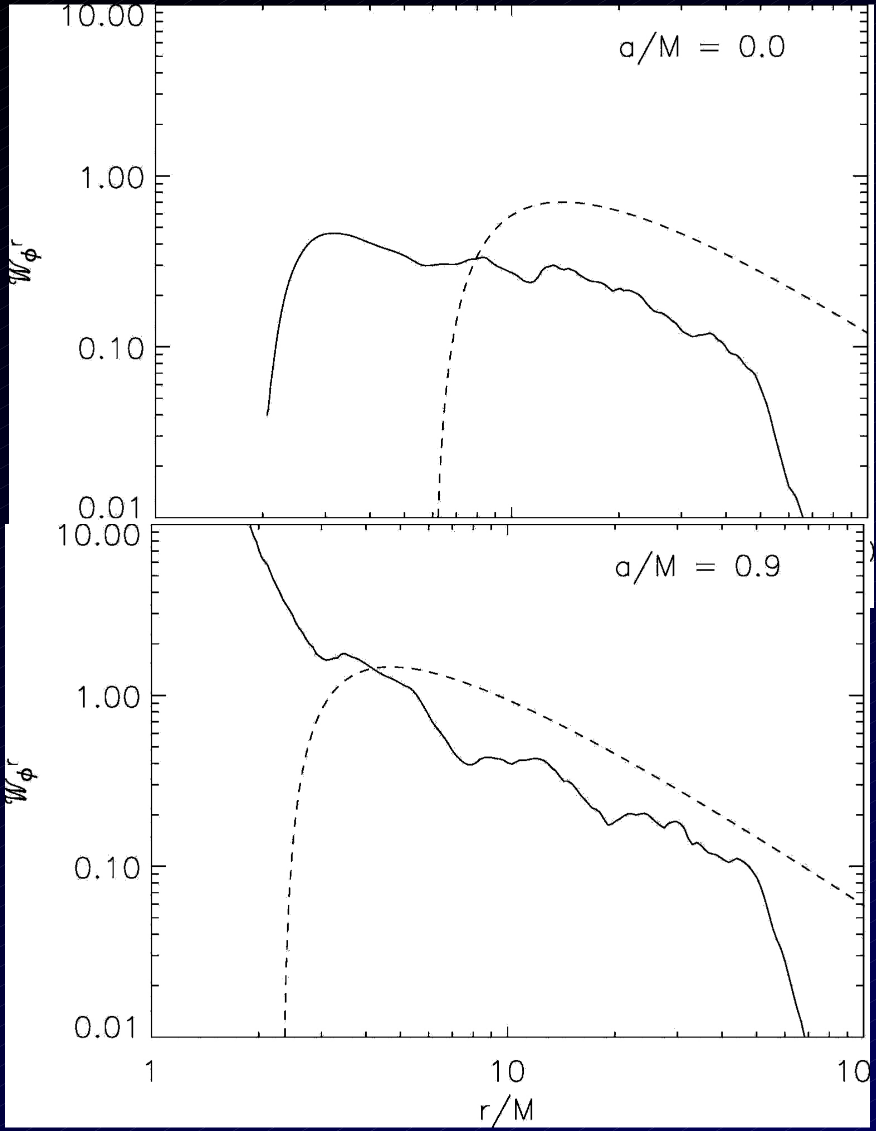
Dynamical Global Disk Models

- De Villiers, Hawley, Hirose, Krolik (2003-2006)
- MRI develops from weak initial field.
- Significant field within ISCO up to the horizon.



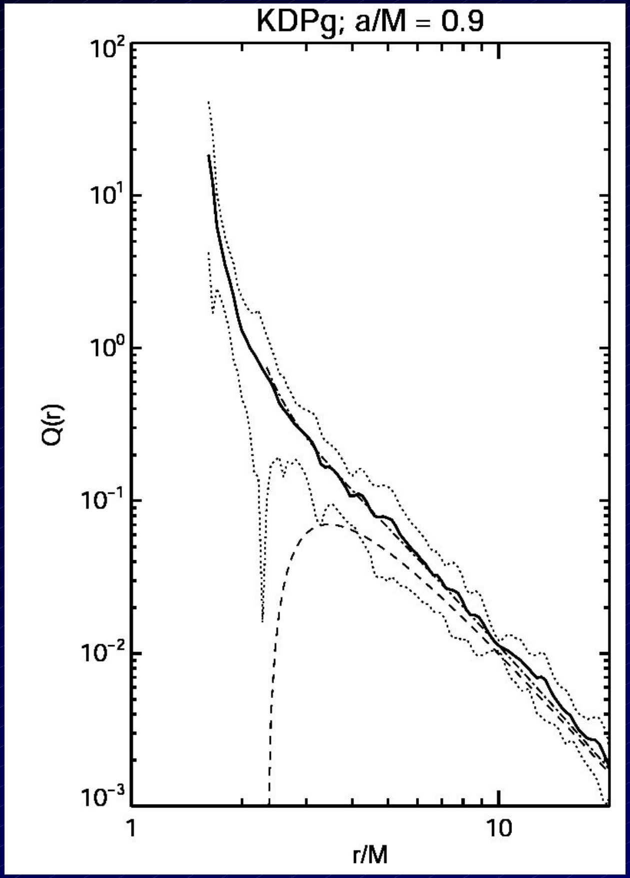
Hirose, Krolik, De Villiers, Hawley (2004)

Dynamical Global Disk Models



Krolik, Hawley, Hirose (2005)

H/R ~ 0.1-0.15



Beckwith, Hawley & Krolik (2008)

- Models dissipation stress as EM stress
- Large dissipation near horizon compensated partially by capture losses and gravitational redshift.
- Used (non-conserv.) int. energy code (dVH) assuming adiabatic flow

Our Method: Simulations with HARM3D

- **HARM**:
  Gammie, McKinney, Toth (2003)

$$\nabla_\nu {}^*F^{\mu\nu} = 0$$

- Axisymmetric (2D)

$$\nabla_\mu \left(\rho u^\mu\right) = 0$$

- Total energy conserving
  (dissipation $\rightarrow$ heat)

- Modern Shock Capturing techniques
  (greater accuracy)

$$\nabla_\mu T^\mu{}_\nu = 0$$

$$T^\mu{}_\nu = \left(\rho + u + p + b^2\right) u^\mu u_\nu + \left(p + \frac{b^2}{2}\right) \delta^\mu{}_\nu - b^\mu b_\nu$$

- Improvements in HARM3D:
  - 3D
  - More accurate
    (parabolic interpolation in reconstruction and constraint transport)
  - Assume flow is isentropic when $P_{gas} << P_{mag}$

SCN, Krolik, Hawley (2009)

Our Method: Simulations with HARM3D

- Improvements:
  - 3D
  - More accurate (higher effective resolution)
  - Stable low density flows
  - Cooling function:
    - Controls energy loss rate
    - Parameterized by H/R
    - $t_{cool} \sim t_{orb}$
    - Only cool when $T > T_{target}$
    - Passive radiation
    - Radiative flux is stored for self-consistent post-simulation radiative transfer calculation

$$\nabla_\nu {}^*F^{\mu\nu} = 0$$

$$\nabla_\mu \left(\rho u^\mu\right) = 0$$

$$\nabla_\mu T^\mu{}_\nu = -\mathcal{F}_\mu$$

$$T(r) = \left(\frac{H}{R} r \Omega\right)^2$$

$$T^\mu{}_\nu = \left(\rho + u + p + b^2\right) u^\mu u_\nu + \left(p + \frac{b^2}{2}\right) \delta^\mu{}_\nu - b^\mu b_\nu$$

SCN, Krolik, Hawley (2009)

GRMHD Disk Simulations

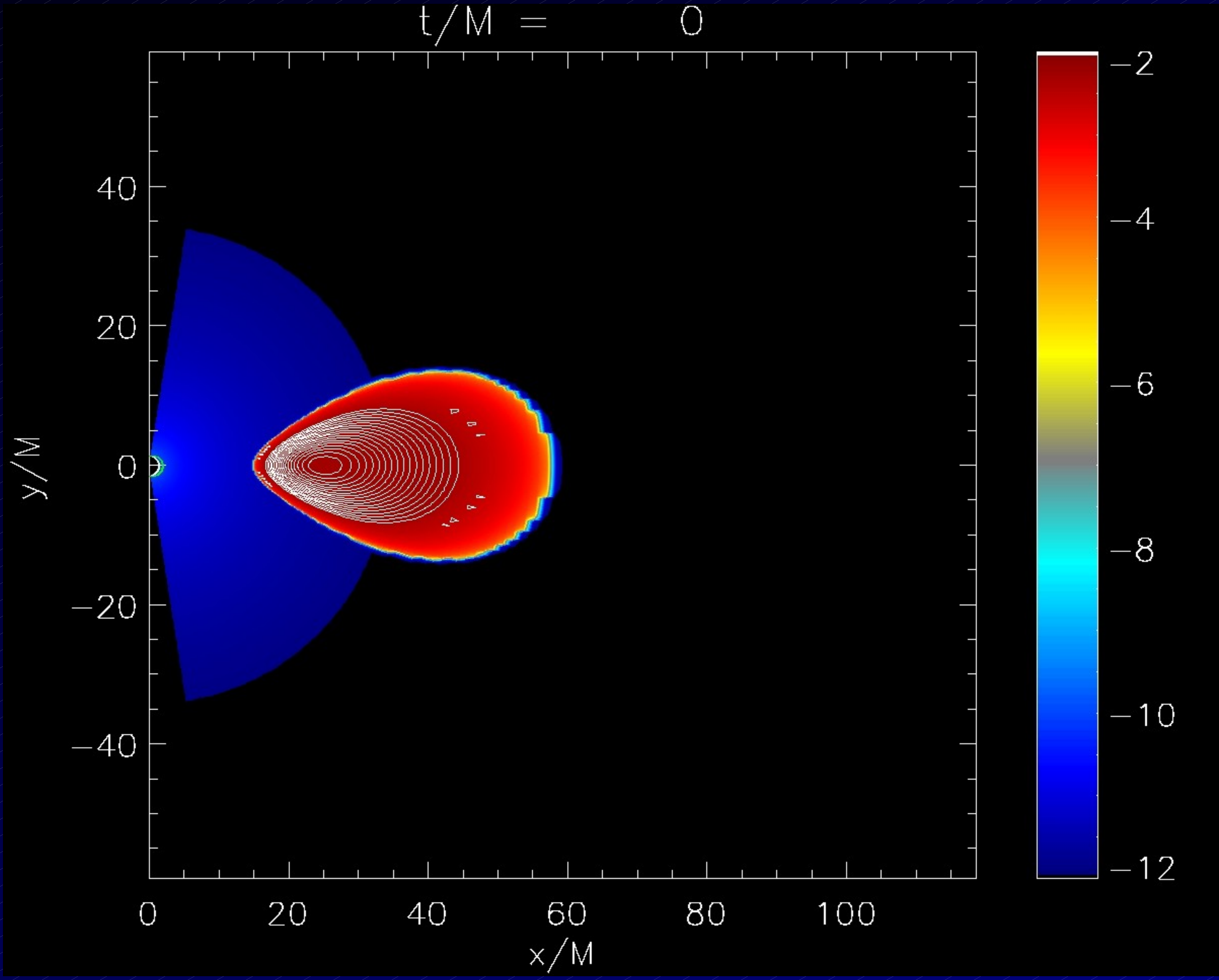
$$N_r \times N_\theta \times N_\phi$$
$$=$$
$$192 \times 192 \times 64$$

$$r \in [r_{hor}, 120\mathrm{M}]$$

$$\theta \in \pi[0.05, 0.95]$$

$$\phi \in [0, \frac{\pi}{2}]$$

$$a = 0.9\mathrm{M}$$



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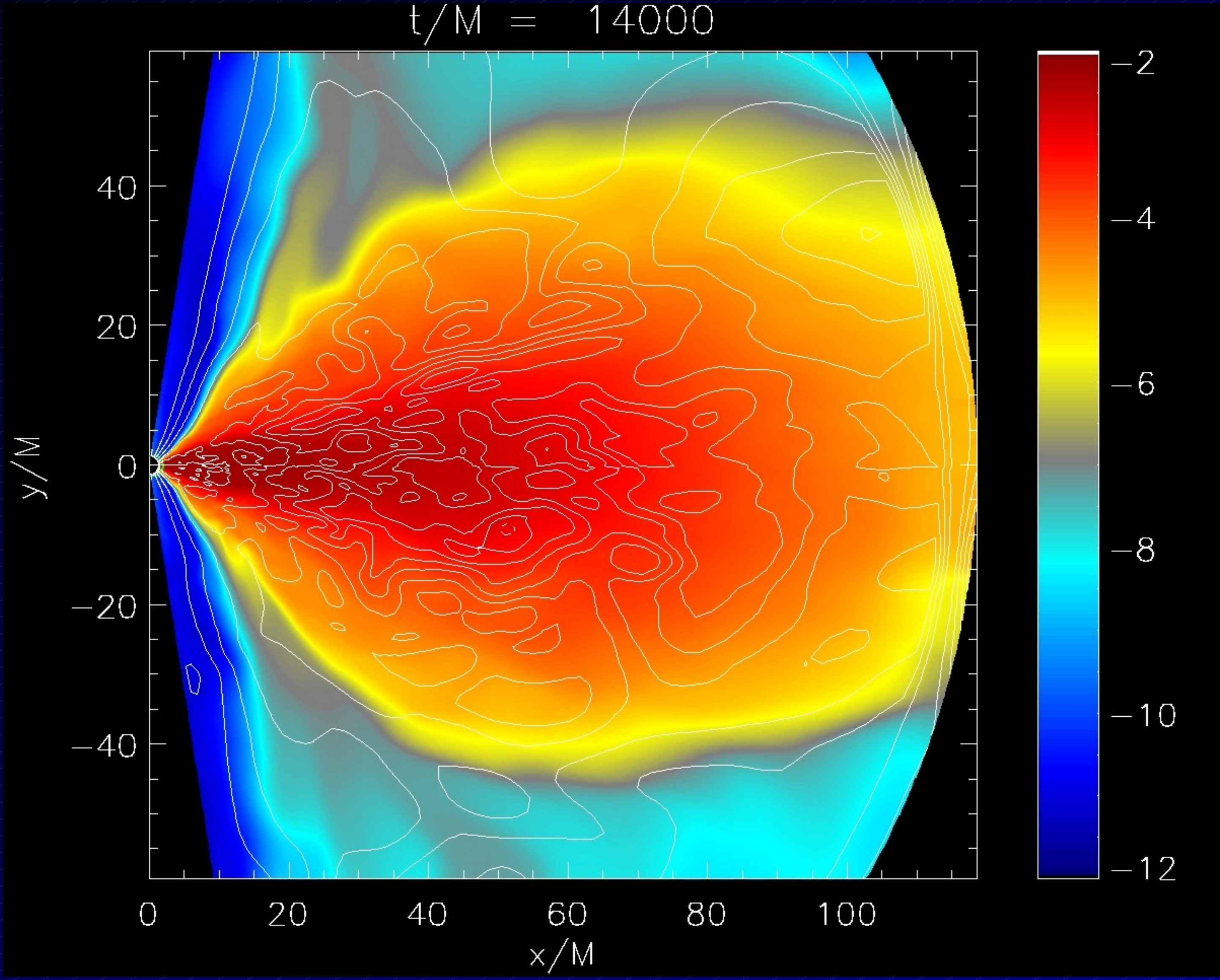
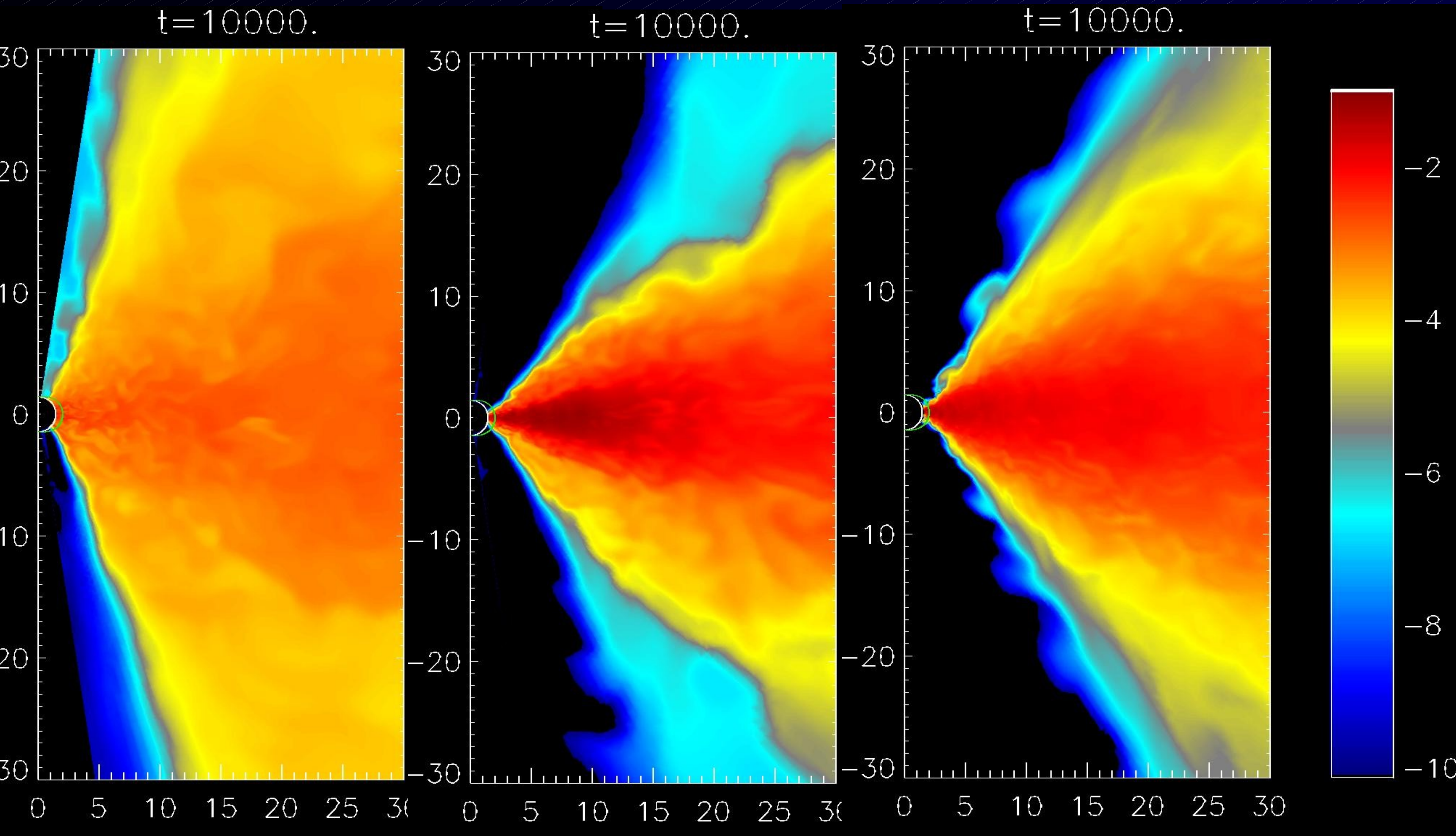
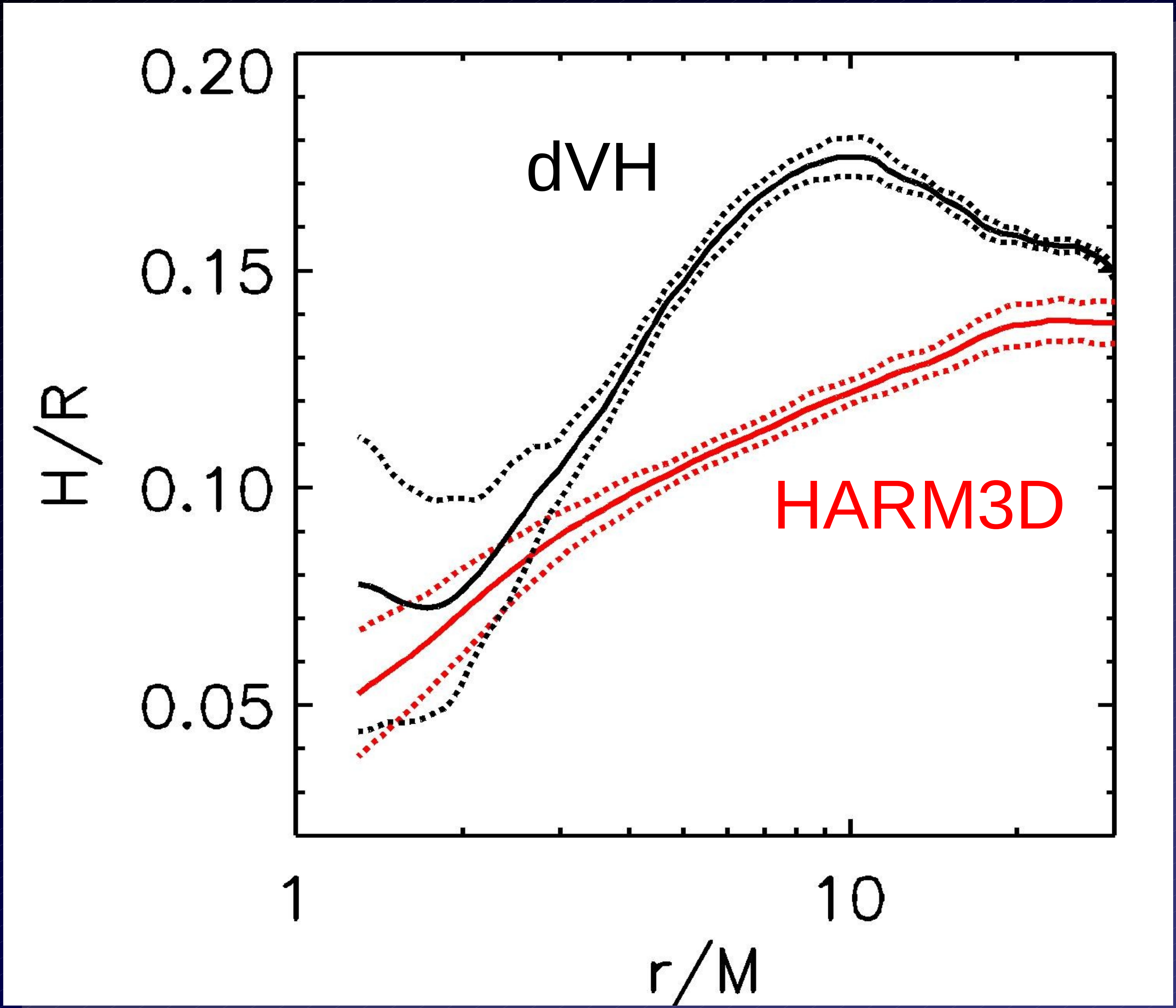


HARM3D vs. dVH $\log(\rho)$



Uncooled **Cooled** **dVH**

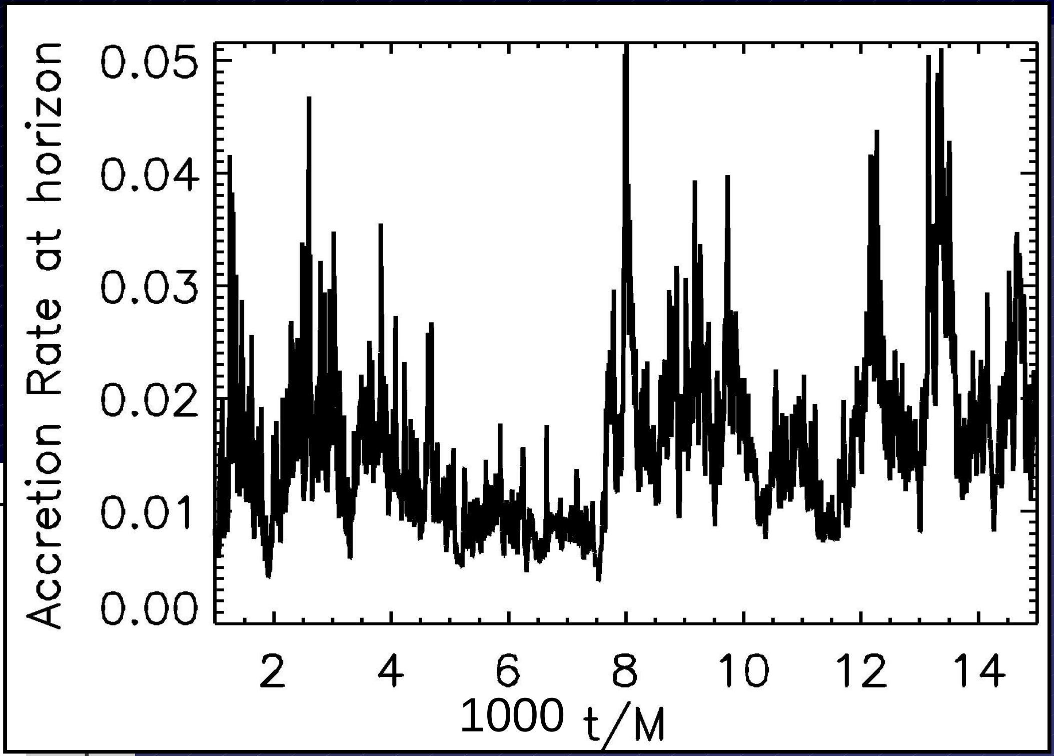
Disk Thickness

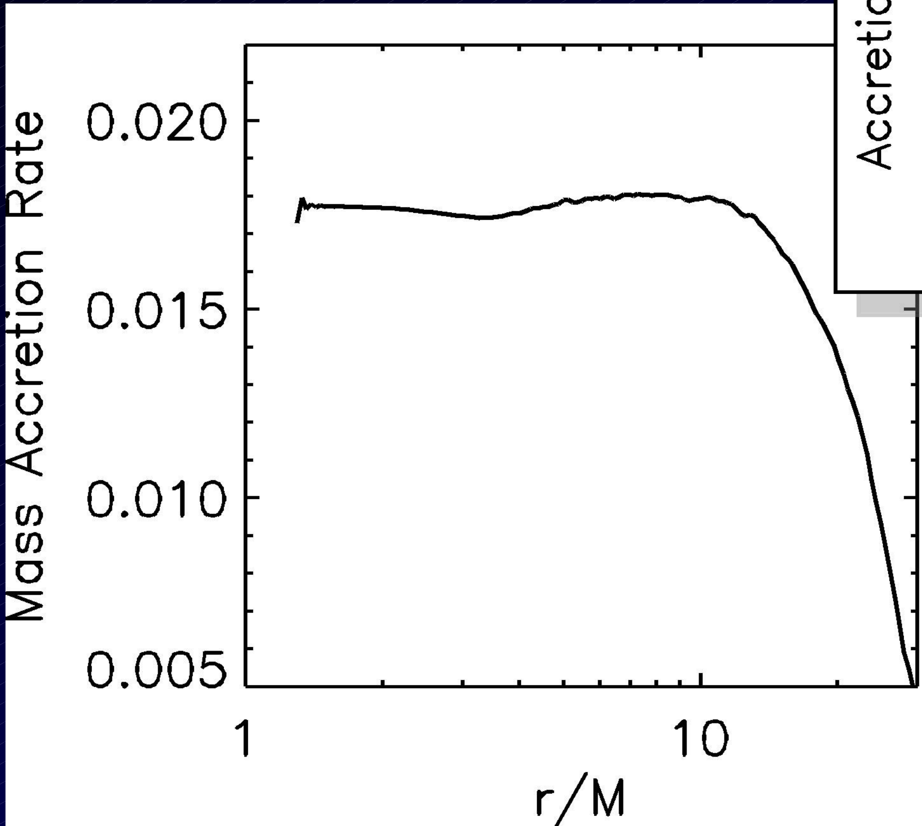


Accretion Rate

Steady State Period = 7000 – 15000M

Steady State Region = Horizon → 12M



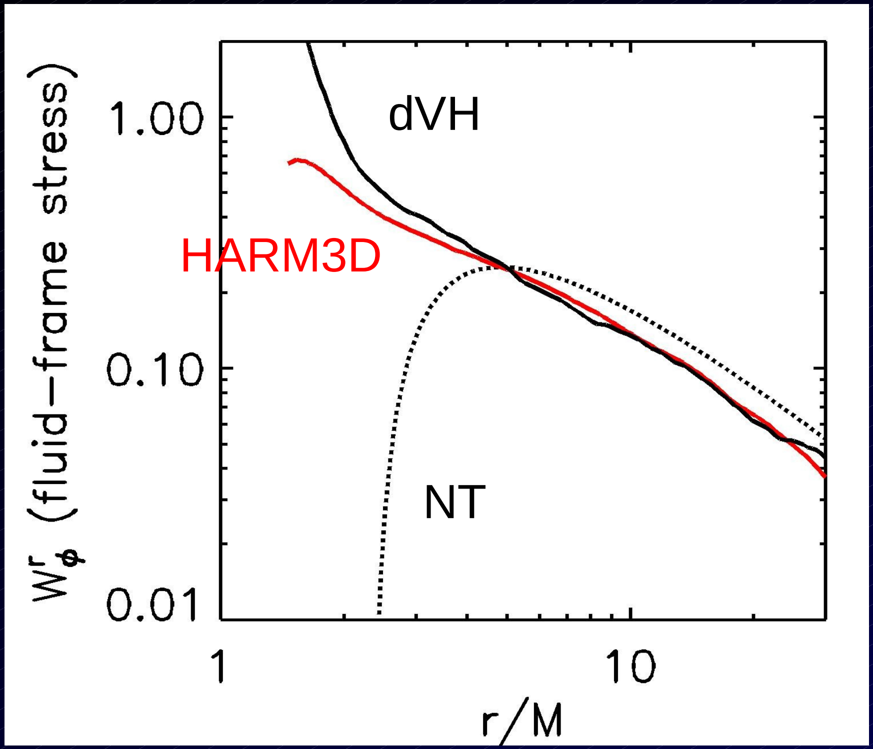


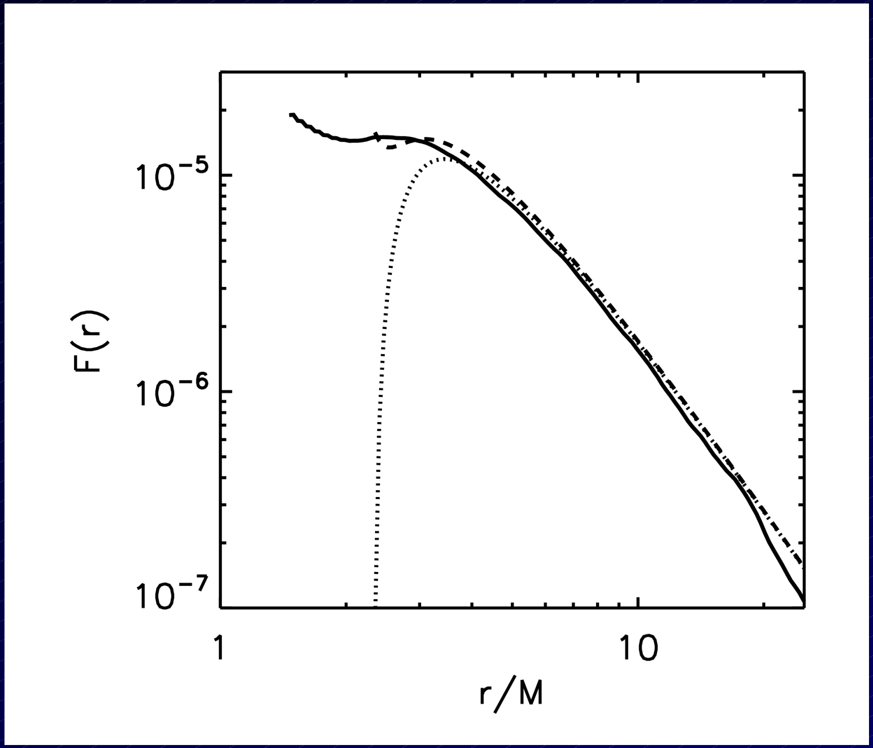
Steady State Important because...

- Match to time-indep. solns
- Real disks are often steady (esp. AGN)
- Provides a baseline for more exotic behavior
- Only in 3D (anti-dynamo in 2D)

$M(r < r_i)$ $\qquad$ $j(r_{horizon}, t)$

Magnetic Stress





Agol & Krolik (2000) model

- Retained Heat → Stress Deficit
- Stress Continuity through ISCO

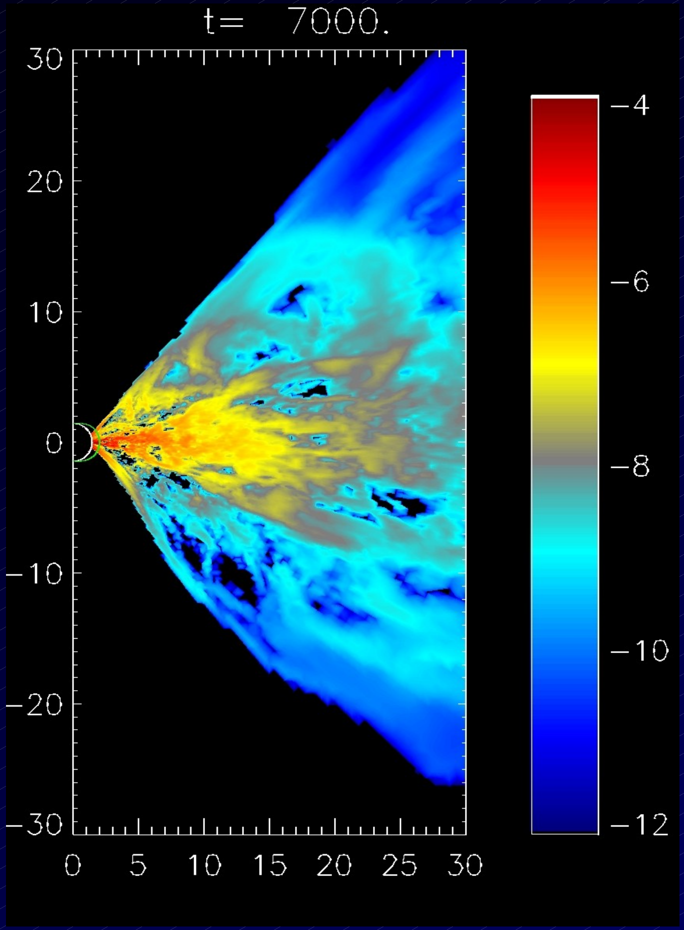
$$\Delta\eta = 0.01$$

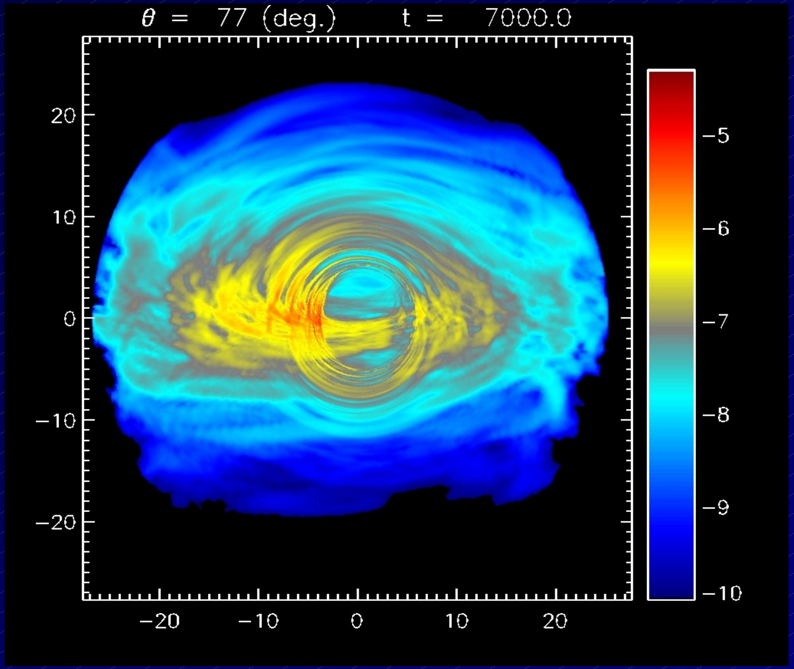
$$\Delta\eta/\eta = 7\ \%$$

Our Method: Radiative Transfer

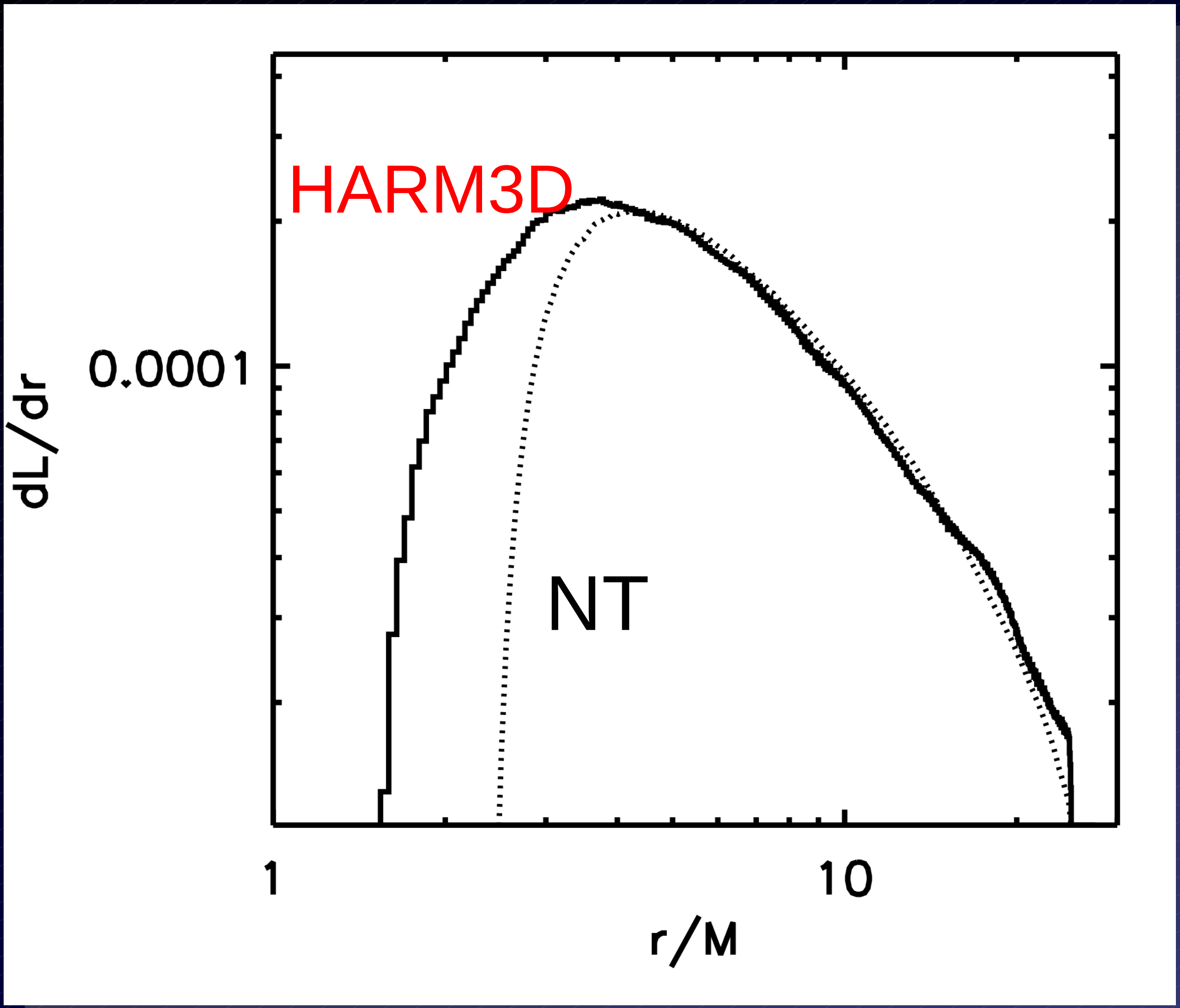
$$j_\nu = \frac{f_c}{4\pi\nu^2}$$

- Full GR radiative transfer
  - GR geodesic integration
  - Doppler shifts
  - Gravitational redshift
  - Relativistic beaming
  - Uses simulation's fluid vel.
  - Inclination angle survey
  - Time domain survey





Observer Frame Luminosity: Angle/Time Average



Assume NT profile for r > 12M .

$$\eta_{H3D} = 0.151$$

$$\eta_{NT} = 0.143$$

$$\Delta\eta/\eta = 6\ \%$$

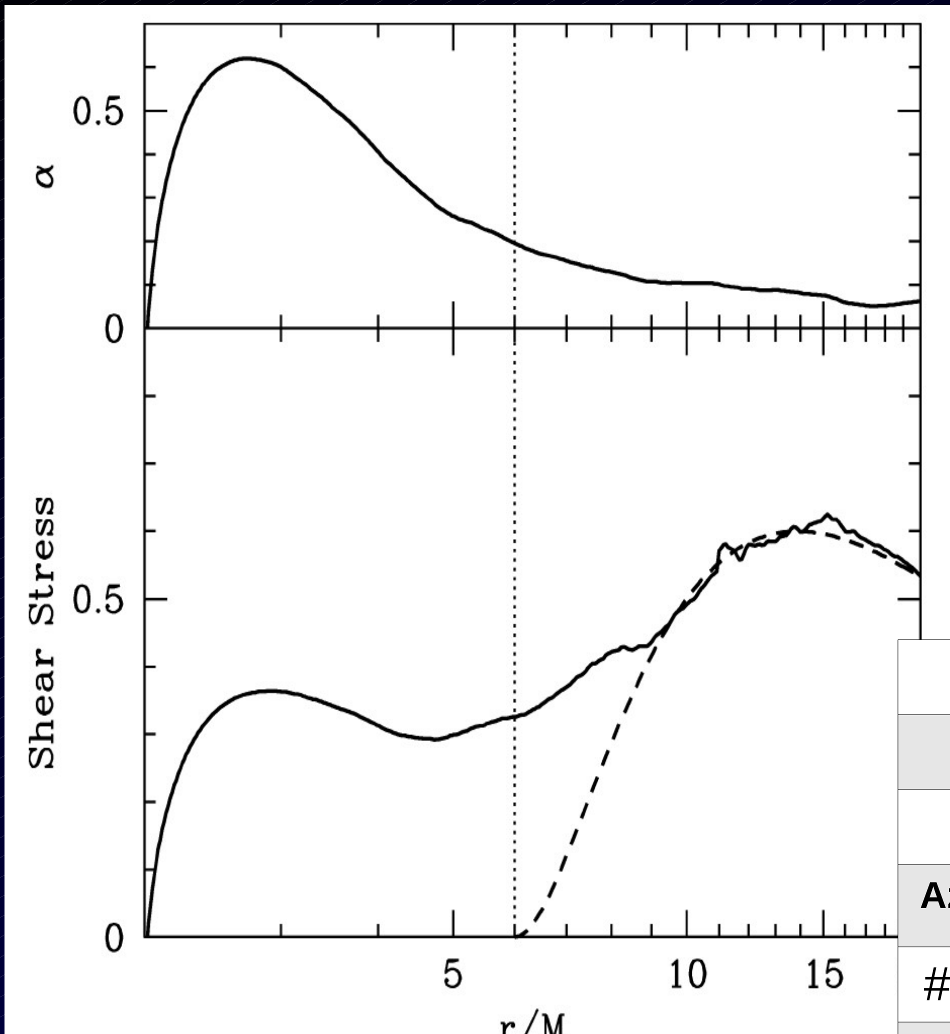
$$\Delta R_{in}/R_{in} \sim 80\%$$

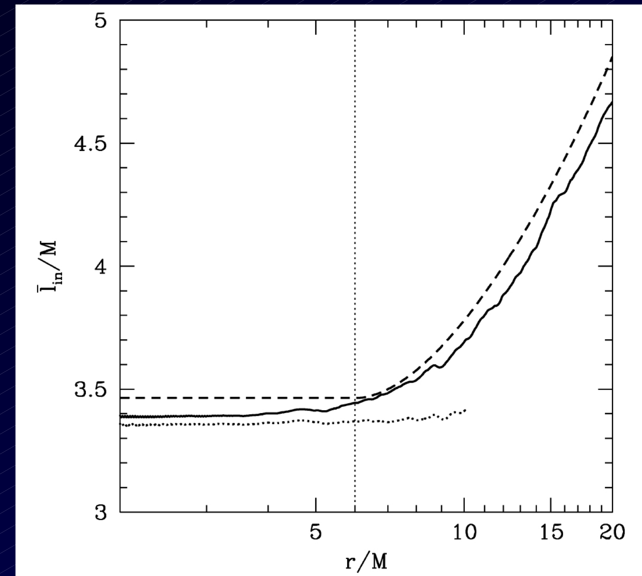
$$\Delta T_{max}/T_{max} = 30\%$$

If disk emitted retained heat: $\Delta\eta/\eta \sim 20\ \%$

SCN, Krolik, Hawley (2009)

Counter Evidence



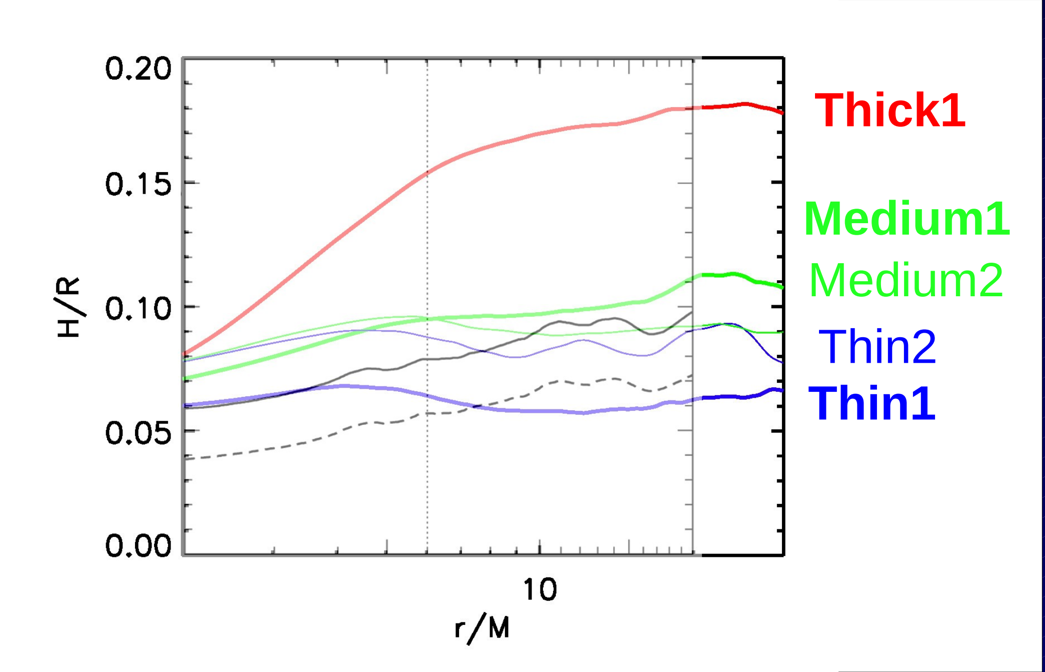


	Shafee et al.	Ours
BH Spin	a=0.0	a=0.9
Resolution	512x120x32	192x192x64
Azimuthal Extent	π/4	π/2
# of B Loops	2	1
H/R	0.05-0.07	0.07-0.13
Code	HARM + 3D	HARM3D

Shafee, McKinney, Narayan, Tchekhovskoy,
Gammie, McClintock (2008)

Counter Counter Evidence

	Theirs	Our Original	Thin1	Medium1	Thick1	Thin2	Medium2
BH Spin	a=0.0	a=0.9	a=0.0	a=0.0	a=0.0	a=0.0	a=0.0
Resolution	512x120x32	192x192x64	912x160x64	512x160x64	384x160x64	192x192x64	192x192x64
ϕ Extent	$\pi/4$	$\pi/2$	$\pi/2$	$\pi/2$	$\pi/2$	$\pi/2$	$\pi/2$
# of Loops	2	1	1	1	1	1	1
Actual H/R	0.05 - 0.07	0.07 - 0.13	0.06	0.10	~0.17	0.087	0.097
N_{cells} per H/r	~60	15 - 30	80	100	40 - 70	60	35
Initial Data	"V. 1"	V. 2	V. 1	V. 1	V. 1	V. 2	V. 2



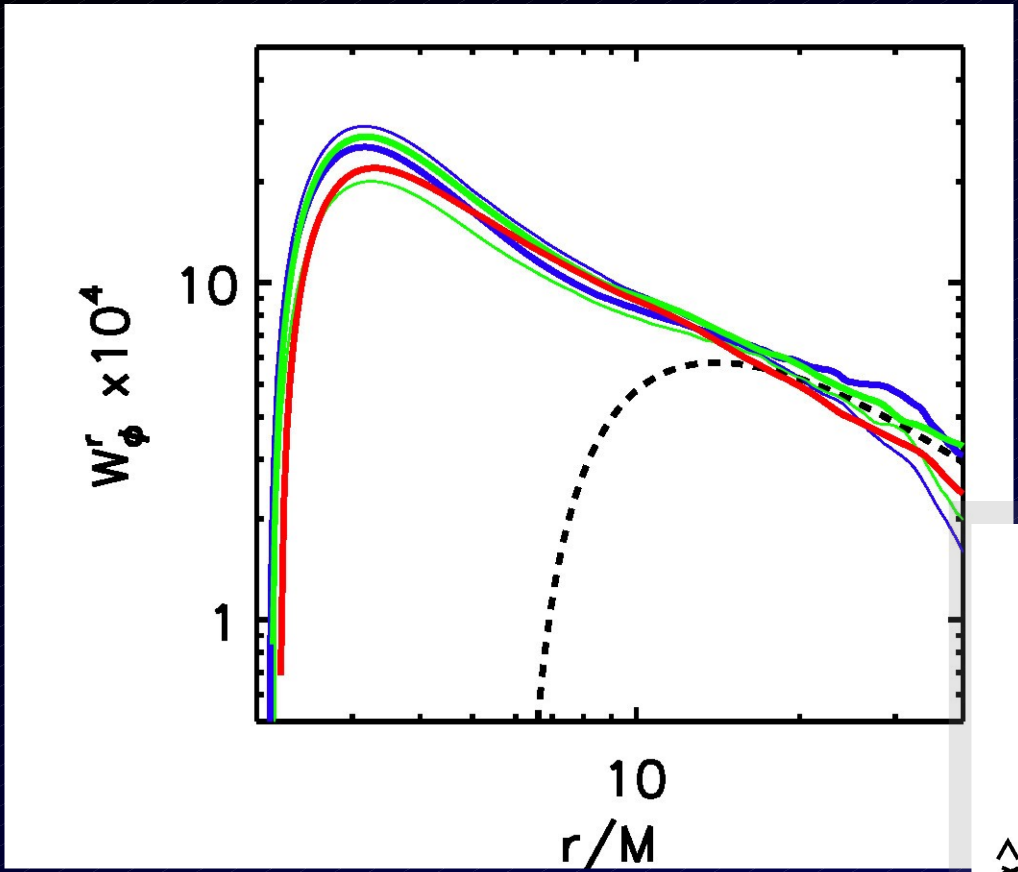
V.1 : Initial disk starts:

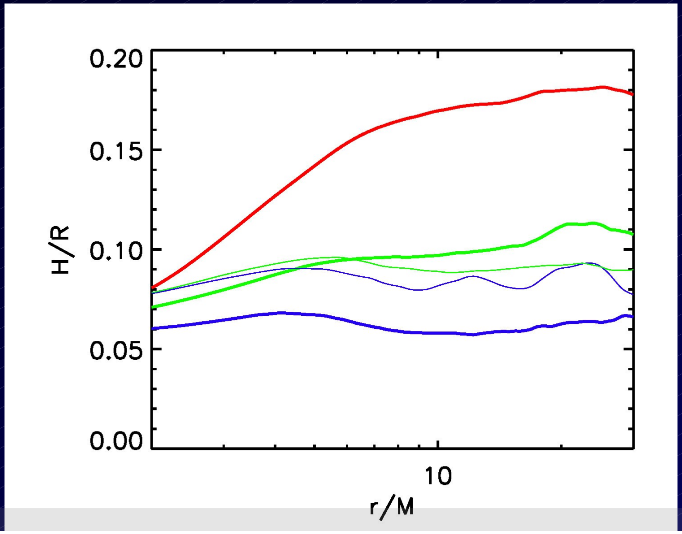
- At target thickness
- With inner radius = 20M
- With p_{max} at r = 35M

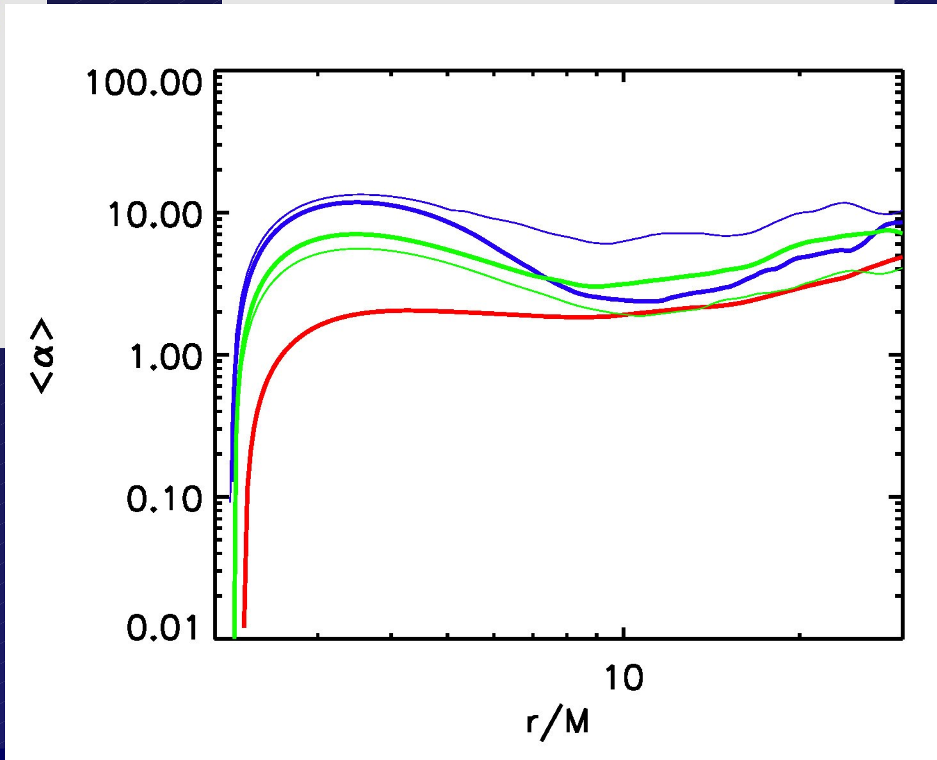
V.2 : Initial disk starts

- At H/R ~ 0.15
- With inner radius = 15M
- With p_{max} at r = 25M

Trends in Scaleheight

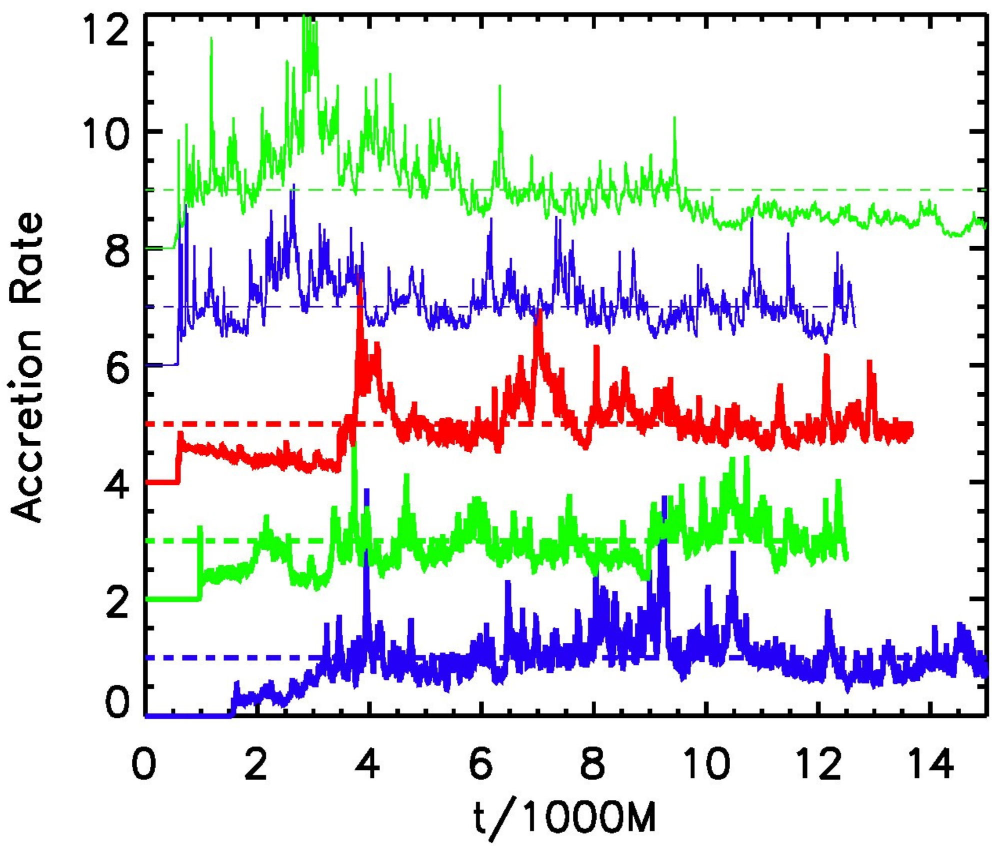


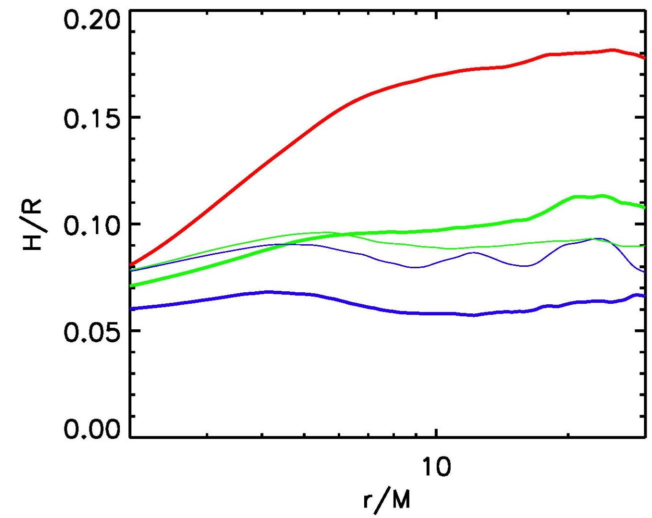


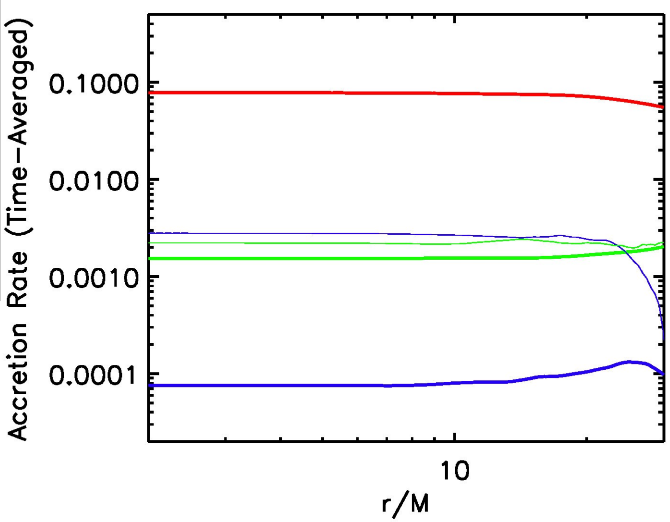


$$W^r_\phi = p\ \alpha$$

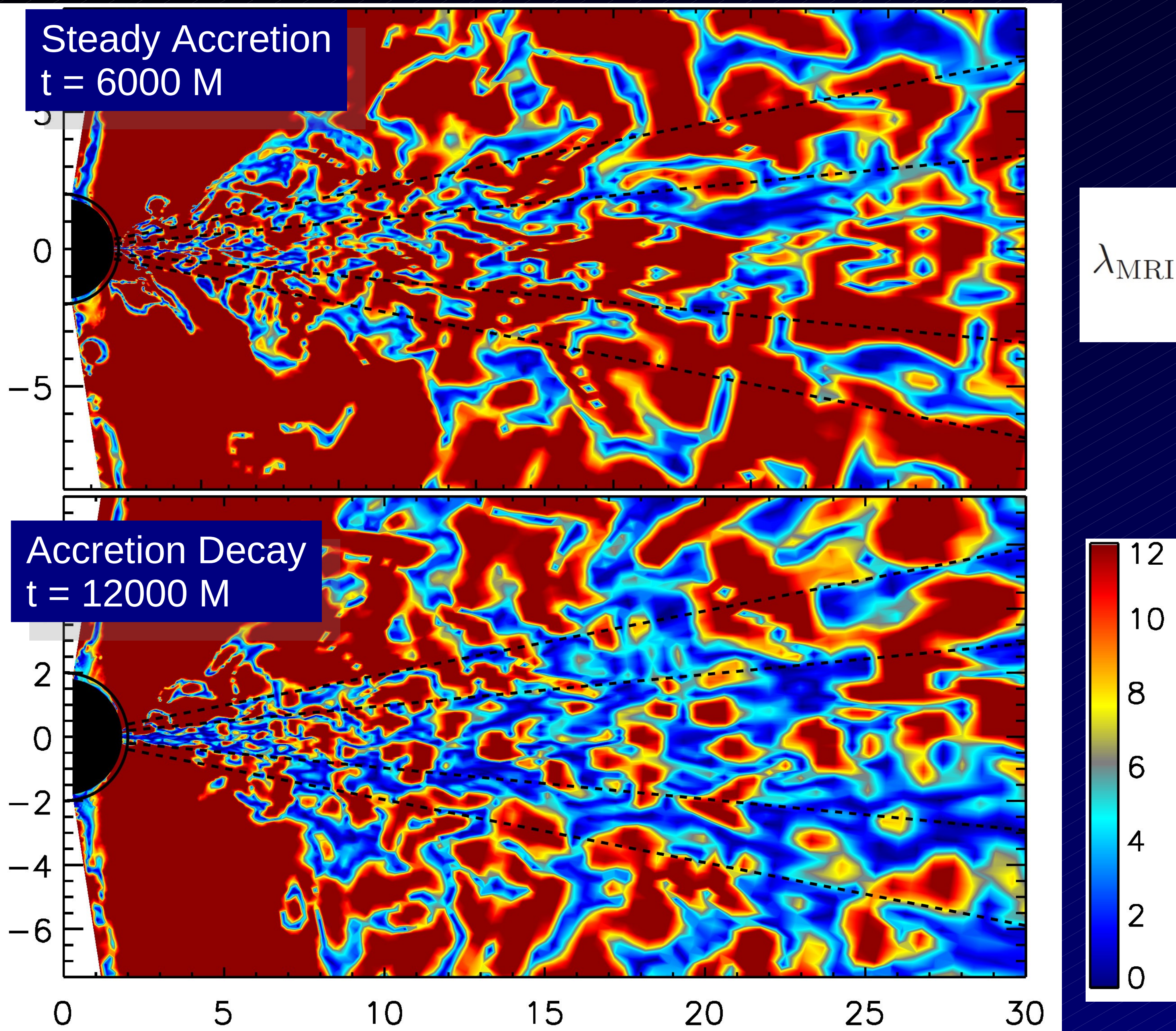
Steady State and Mass Flow Equilibrium







Resolution of the MRI

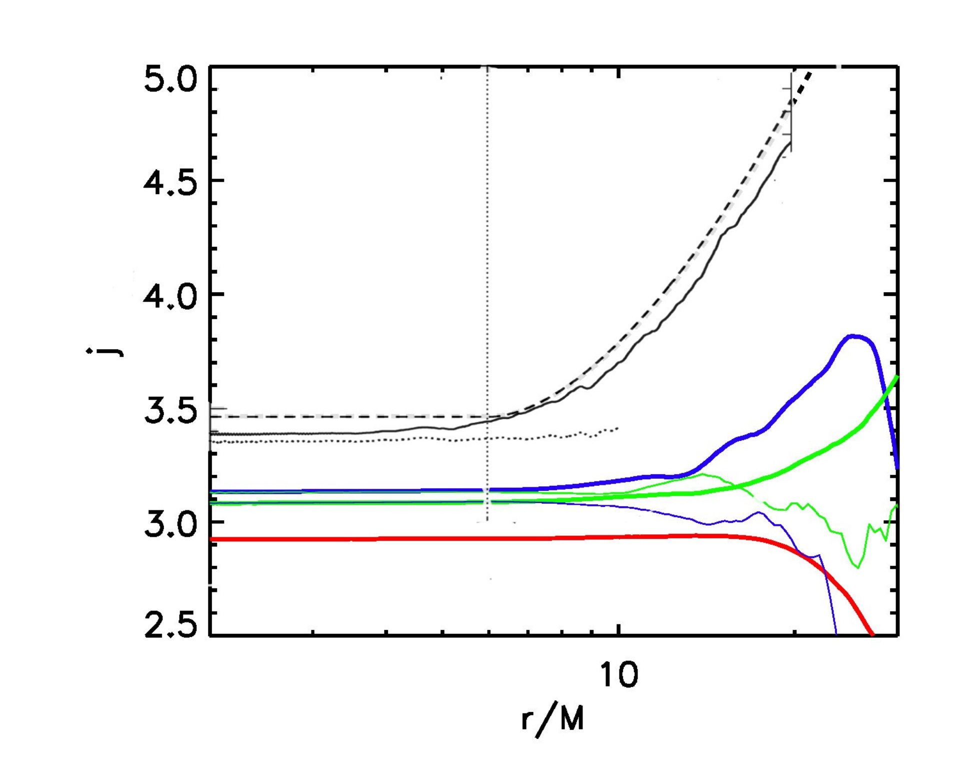


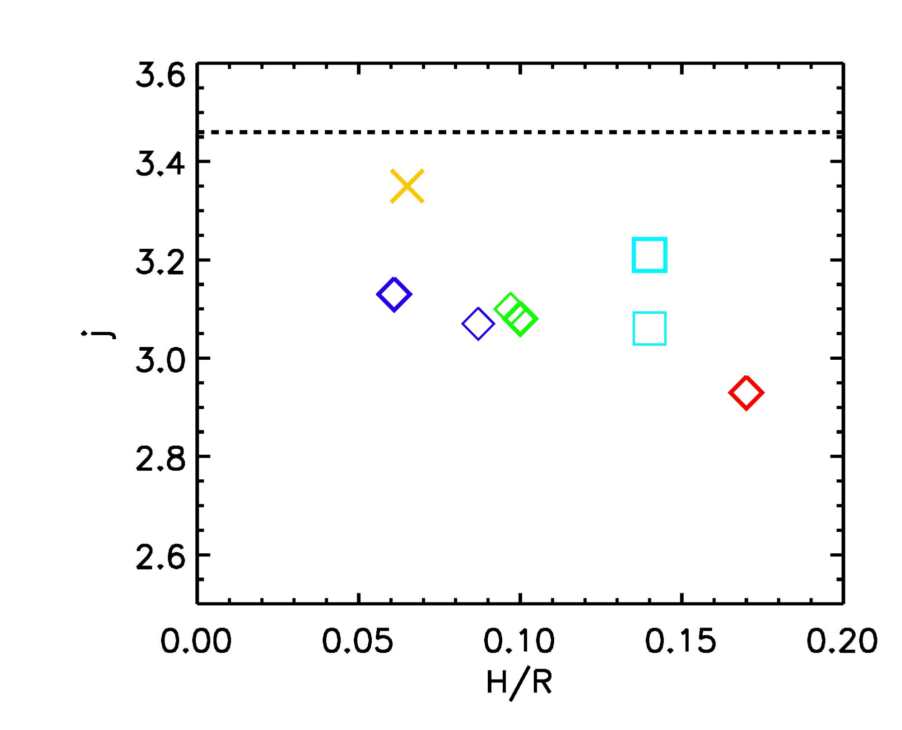
$$\lambda_{\rm MRI} \equiv \frac{1}{\sqrt{4\pi\rho\Omega(R)}} b_\mu \hat{e}^\mu_{(\theta)}$$

$$\frac{\lambda_{\rm MRI}}{\Delta z} > 6$$

Sano et al. (2004)

Accreted Specific Angular Momentum





- Dependence is weak ~ $(H/R)^{(1/2)}$ instead of "expected" $(H/R)^2$
- Possible Dependence on Initial Field Topology
- Independent of Algorithm (modulo Shafee et al. 2008)
- Still need to transport radiated energy to infinity to find efficiency

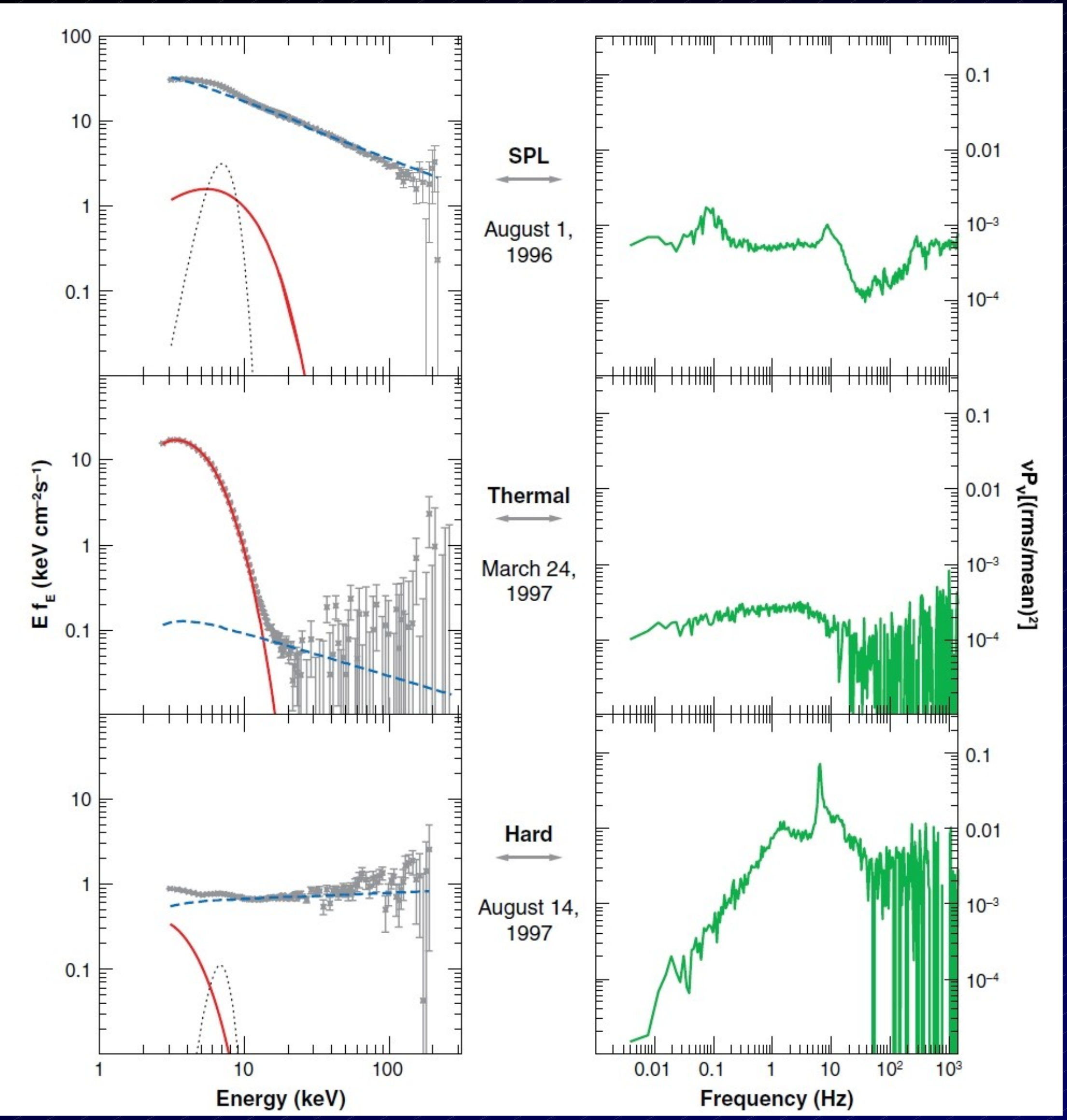
3D HARM w/ 2 Poloidal Loops
Shafee et al. (2008)

HARM3D 1 Poloidal Loop

GRMHD (dVH) 1 Poloidal Loop

GRMHD (dVH) Vertical Field
Beckwith, Hawley, Krolik (2009)

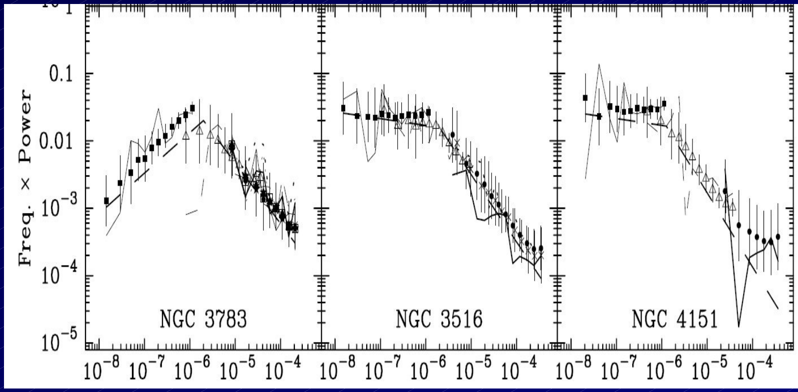
X-ray Variability of Accretion



- X-ray var. always dominated by corona
- XRB var. dependent on spectral state

$$P \sim \nu^{\alpha}$$

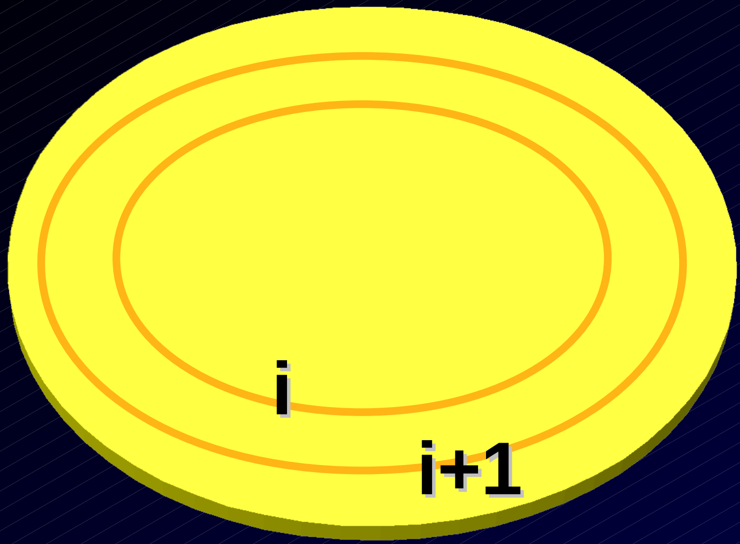
$$-3 < \alpha < -1$$



XRBs: Remillard & McClintock (2006)

AGN: Markowitz et al.(2003)

Variability Models



$$P \sim \nu^{\alpha}$$

Lyubarskii (1997)

Total variability is a superposition of independent variability from larger radii modulating interior annuli on inflow time scales

Churazov, Gilfanov, Revnivtsev (2001)

Outer radius of corona may be cause of (temporal) spectral slope.

$$\tau_{\mathrm{a}} = \left[\alpha \left(\frac{H}{r} \right)^{2} \Omega_{\mathrm{K}} \right]^{-1}$$

- Accretion rate modulation modeled as variability of α

- Predict phase coherence at frequencies longer than inflow freq.

Armitage & Reynolds (2003)
Machida & Matsumoto (2004)
Schnittman et al. (2006)
Reynolds & Miller (2009)

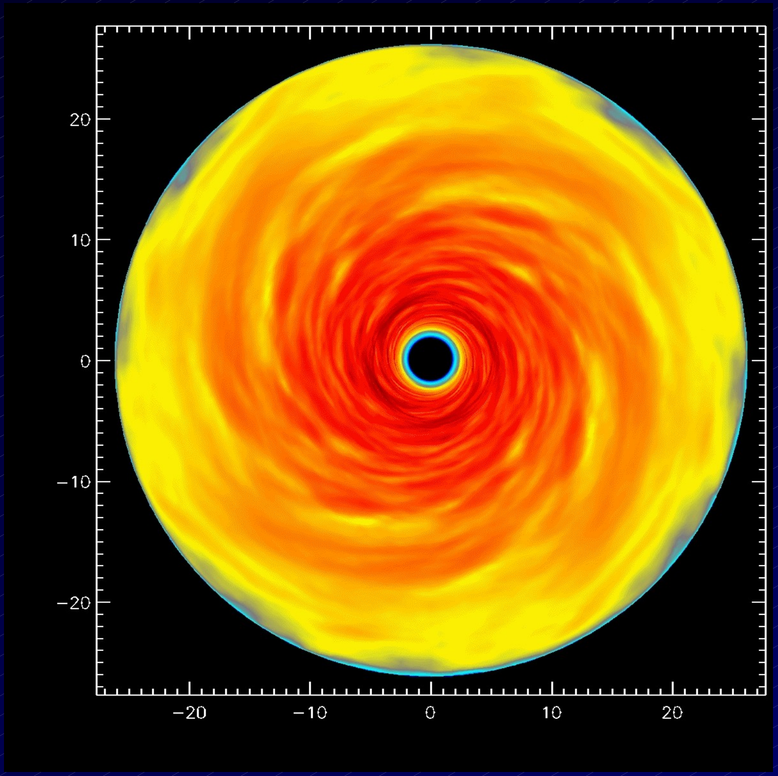
- Used accretion rate or stress as dissipation proxies
- PLD breaks at local orbital frequency per annulus
- $\alpha_{\omega<\Omega} \sim -1 \quad \alpha_{\omega>\Omega} \sim -3$
- Composite PLD $\alpha \sim -2$

Our Variability Model

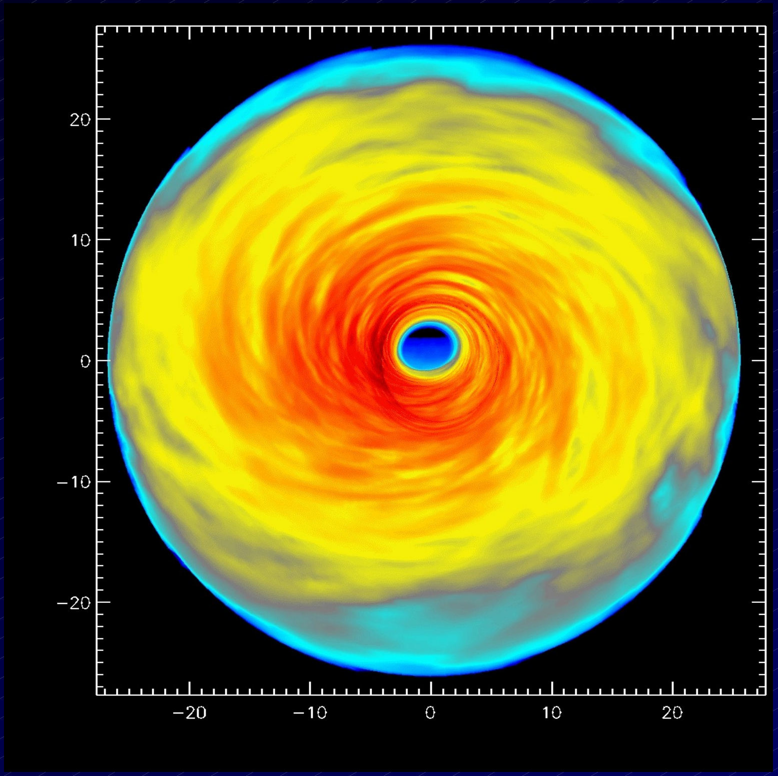
Noble & Krolik (2009)

Simulation: a = 0.9M   H/R = 0.07 - 0.13

- Assume Thomson Scattering
- Optical depth set by $\dot{m} = L/\eta L_E$
- Integrate emission up to photosphere
- Include effect of finite light speed
- Parameterized by $\theta, \dot{m}$

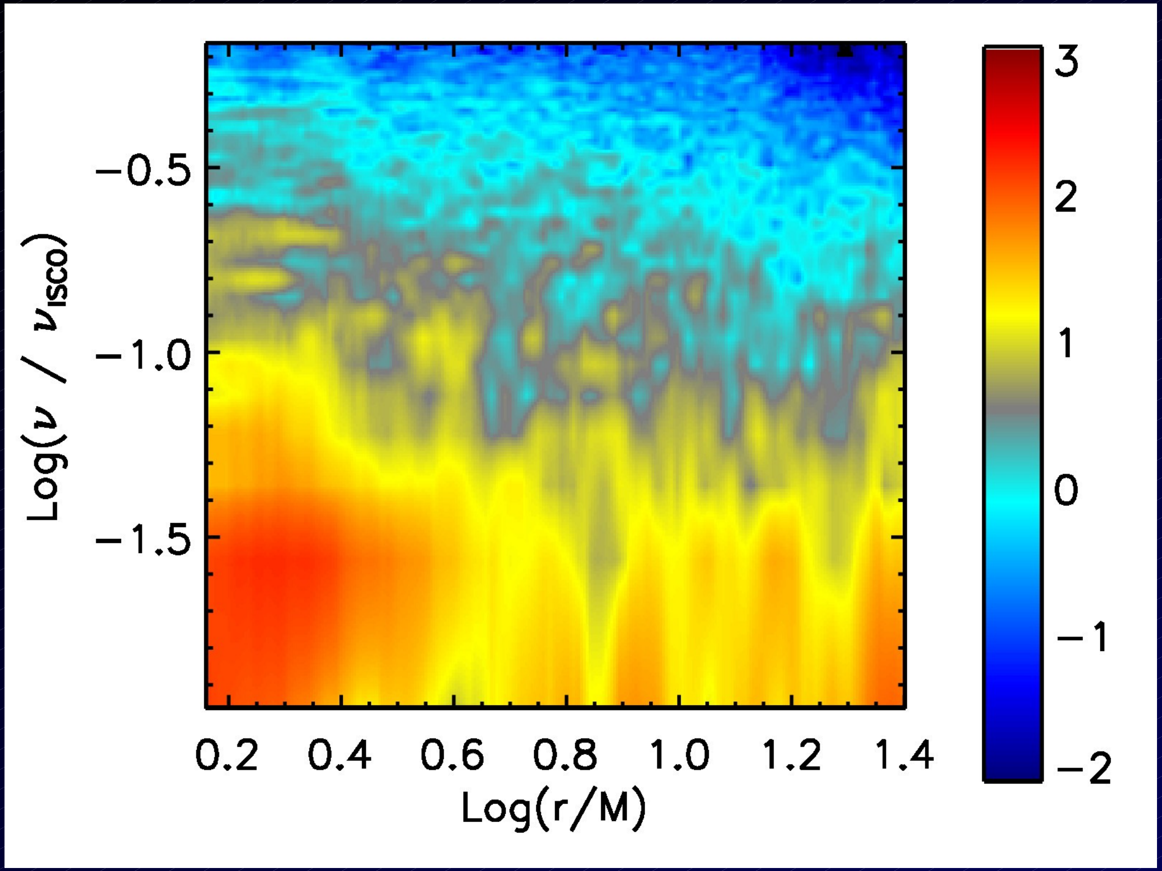


$\dot{m} = 0.003$



$\theta = 41^o$

Spectra of Annuli



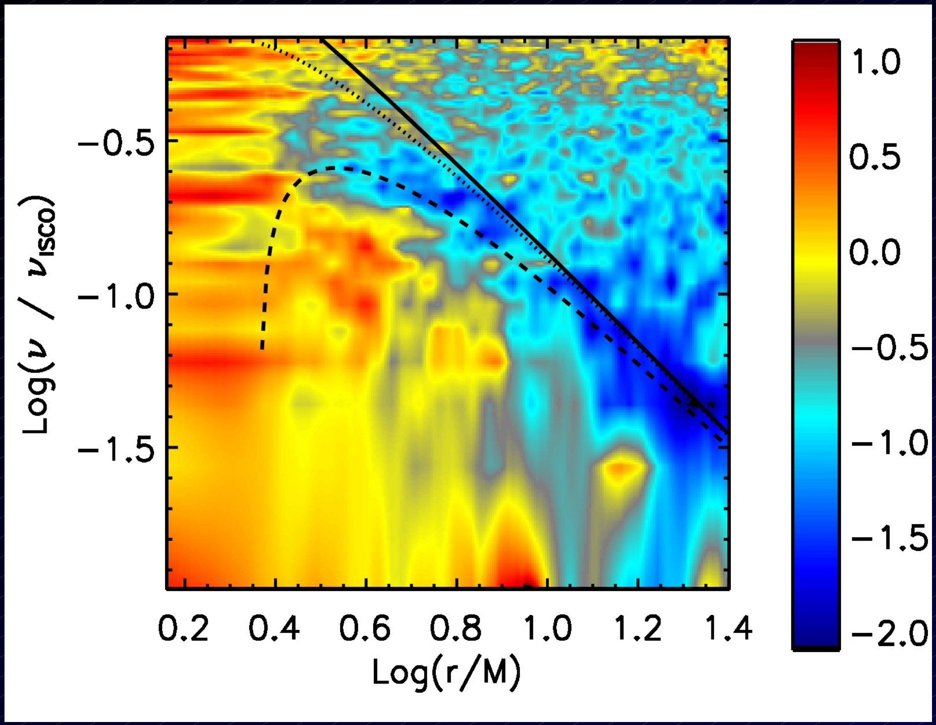
- No PL break at Ω
- Each annulus $\alpha \sim -2$
- More power at smaller r
- No feature at ISCO

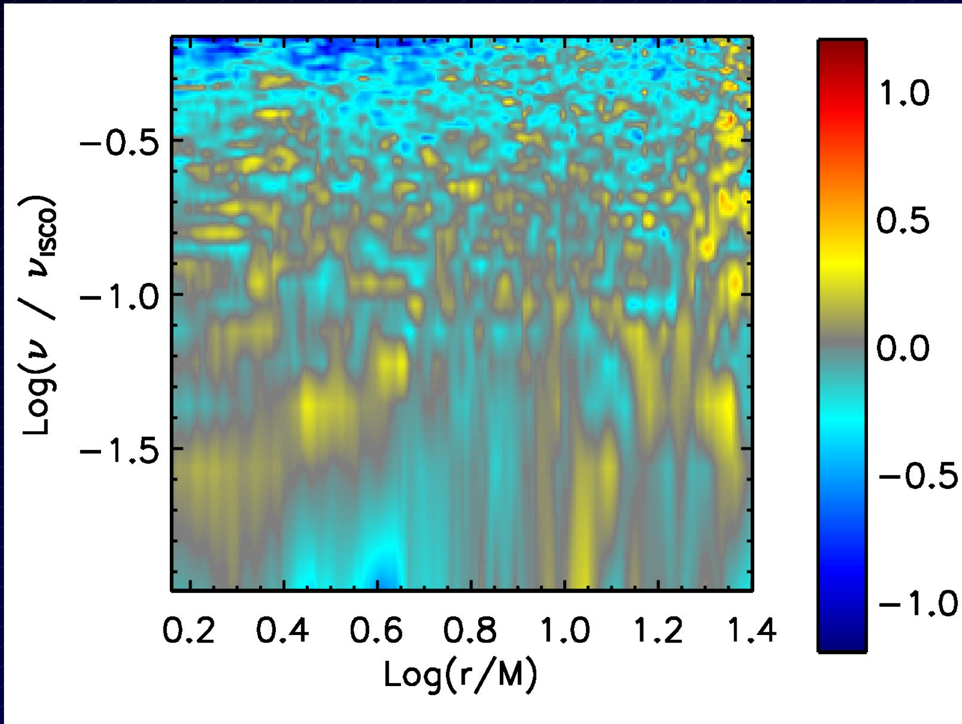
$$F(t,r) = \frac{dL}{dr}(r,t)$$

$$\hat{F}(\nu,r) = \frac{1}{N}\sum_{n=0}^{N-1} F(t_n,r)e^{-2\pi i\nu t_n}$$

$$P(\nu,r) = \frac{2T}{\bar{F}^2}\left|\hat{F}\left(\nu,r\right)\right|^2$$

Origin of Variability





$$P_{diss}(\nu,r)/P_{\dot{M}}(\nu,r)$$

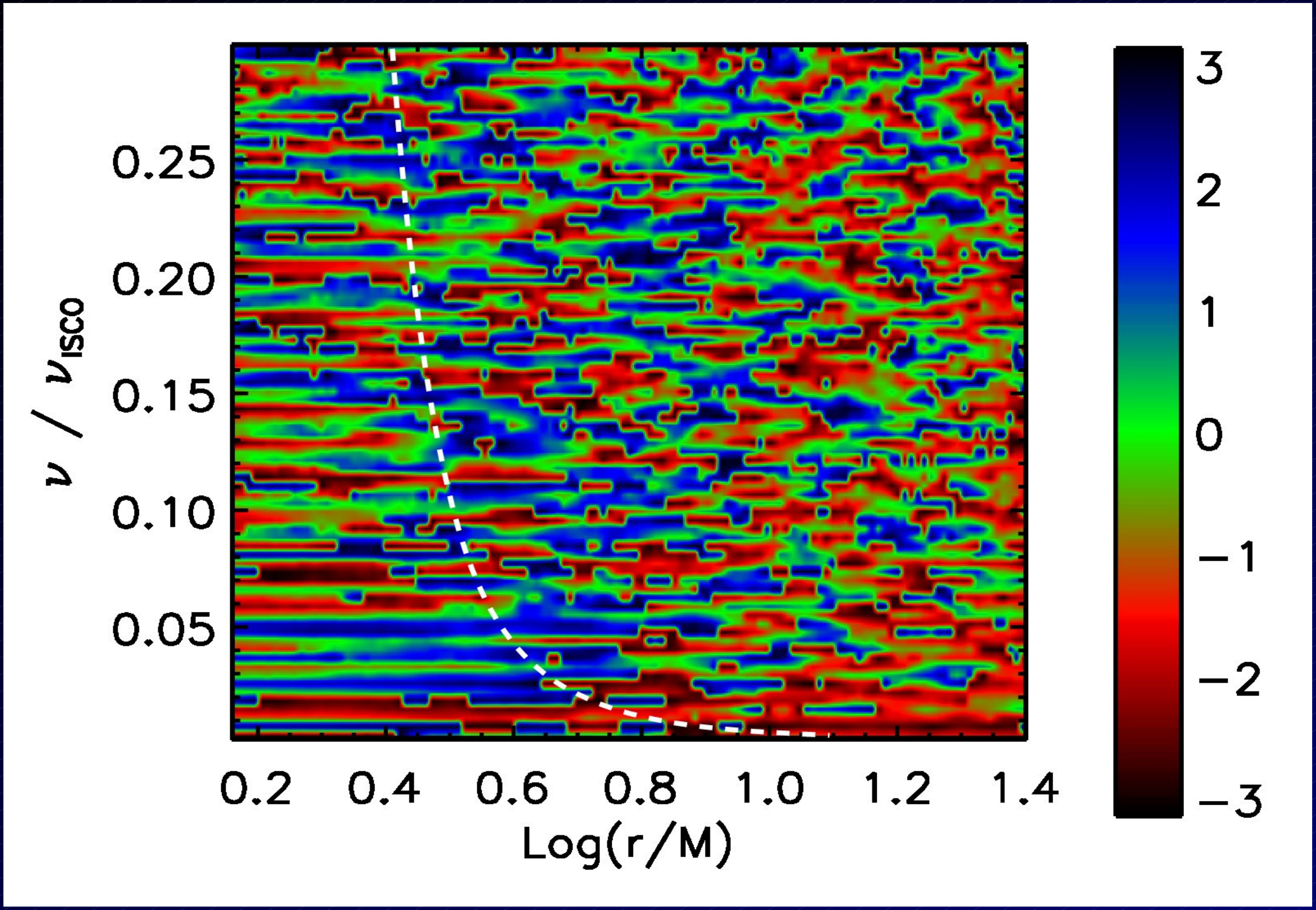
$$P_{inf}(\nu,r)/P_{diss}(\nu,r)$$

$$\theta=5^o$$

- Epicyclic motion not dissipated
- Dissipation not well proxied by $\dot{M}$

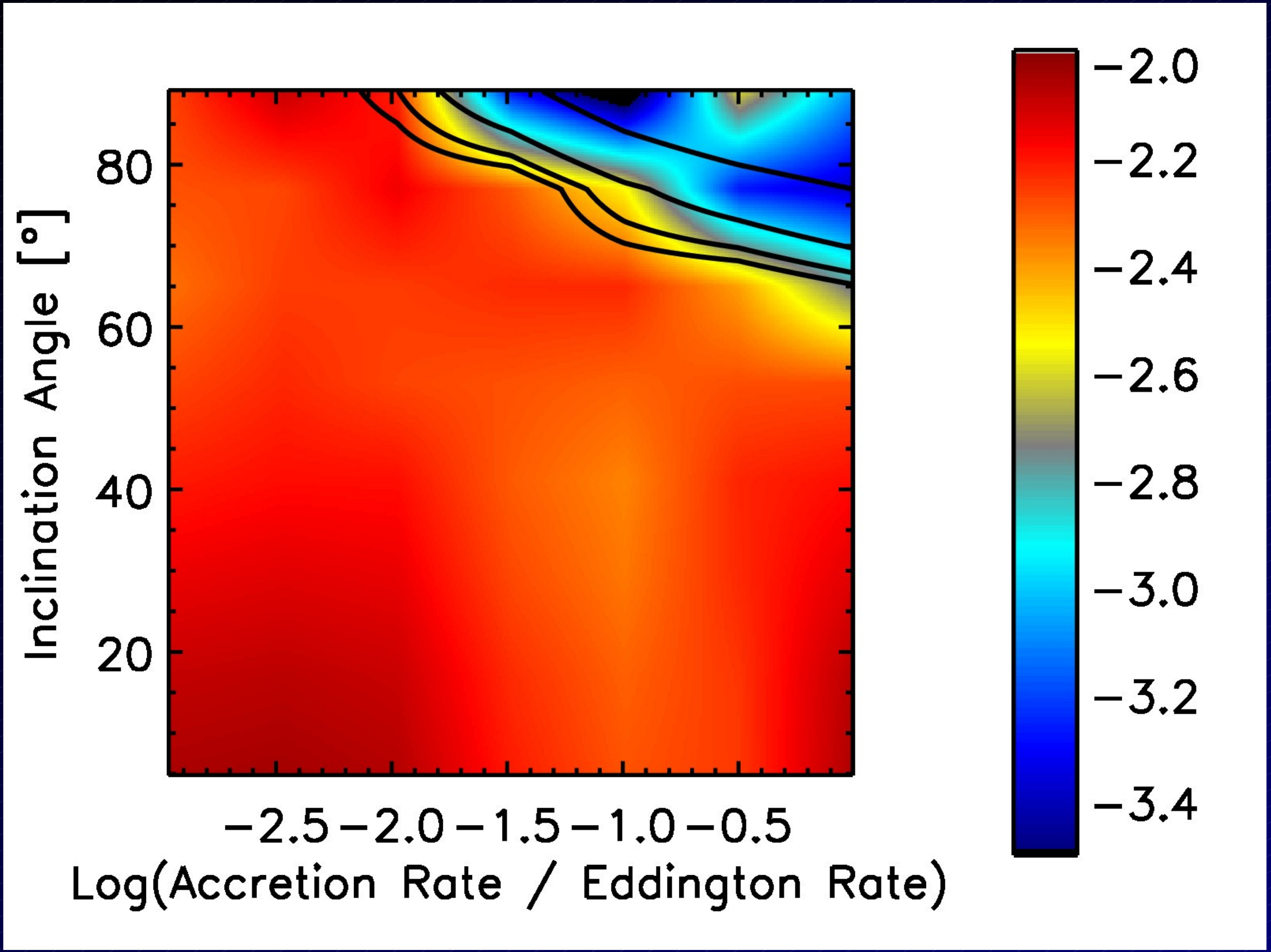
- Observed var. ~ local dissipation var.

Phase Coherence



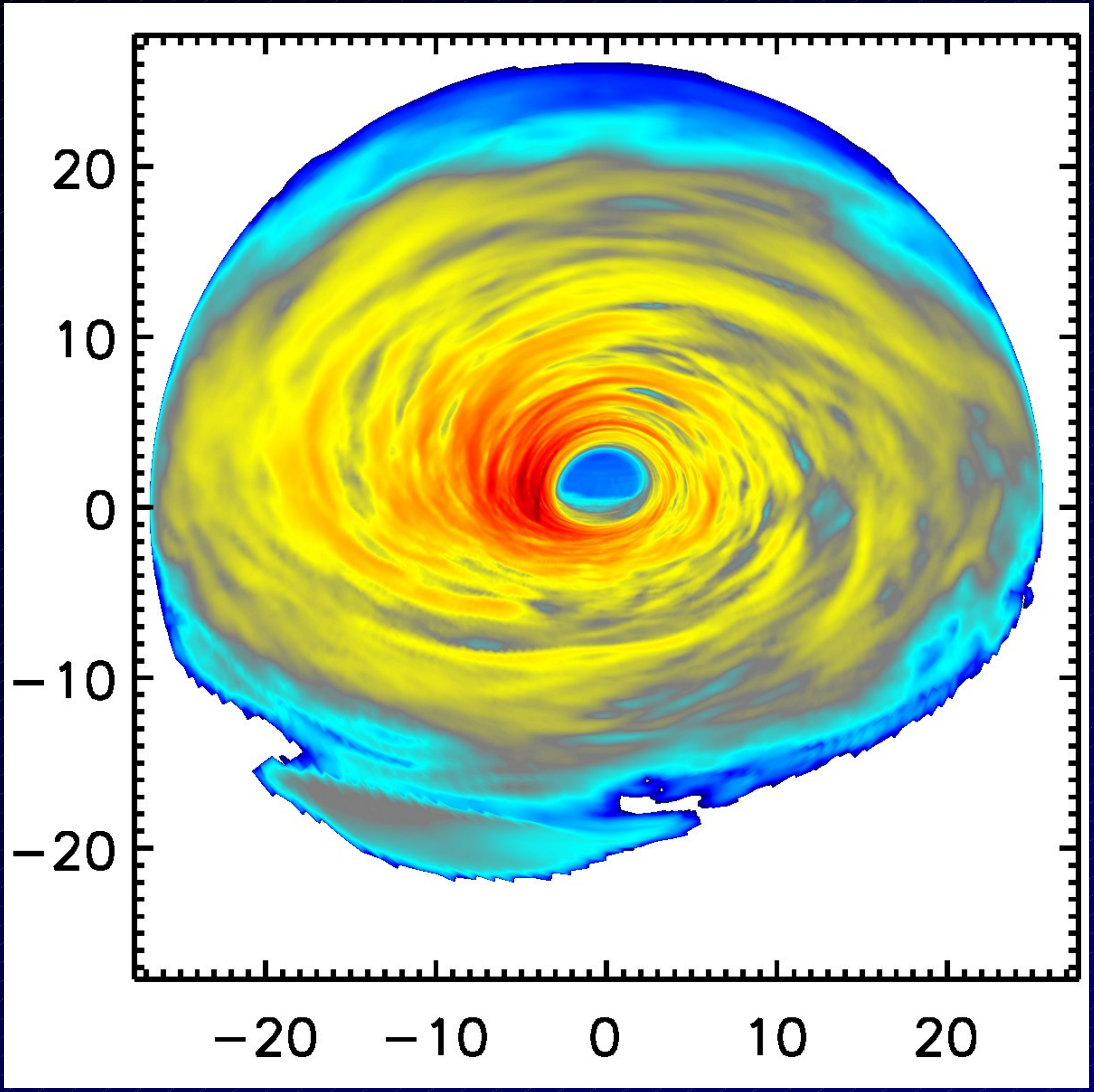
- Possible coherence below inflow frequency (ala Lyubarskii)
- Otherwise dissipation is incoherent over all scales

PLD Exponent vs. Parameter Space



- Complete degeneracy!!

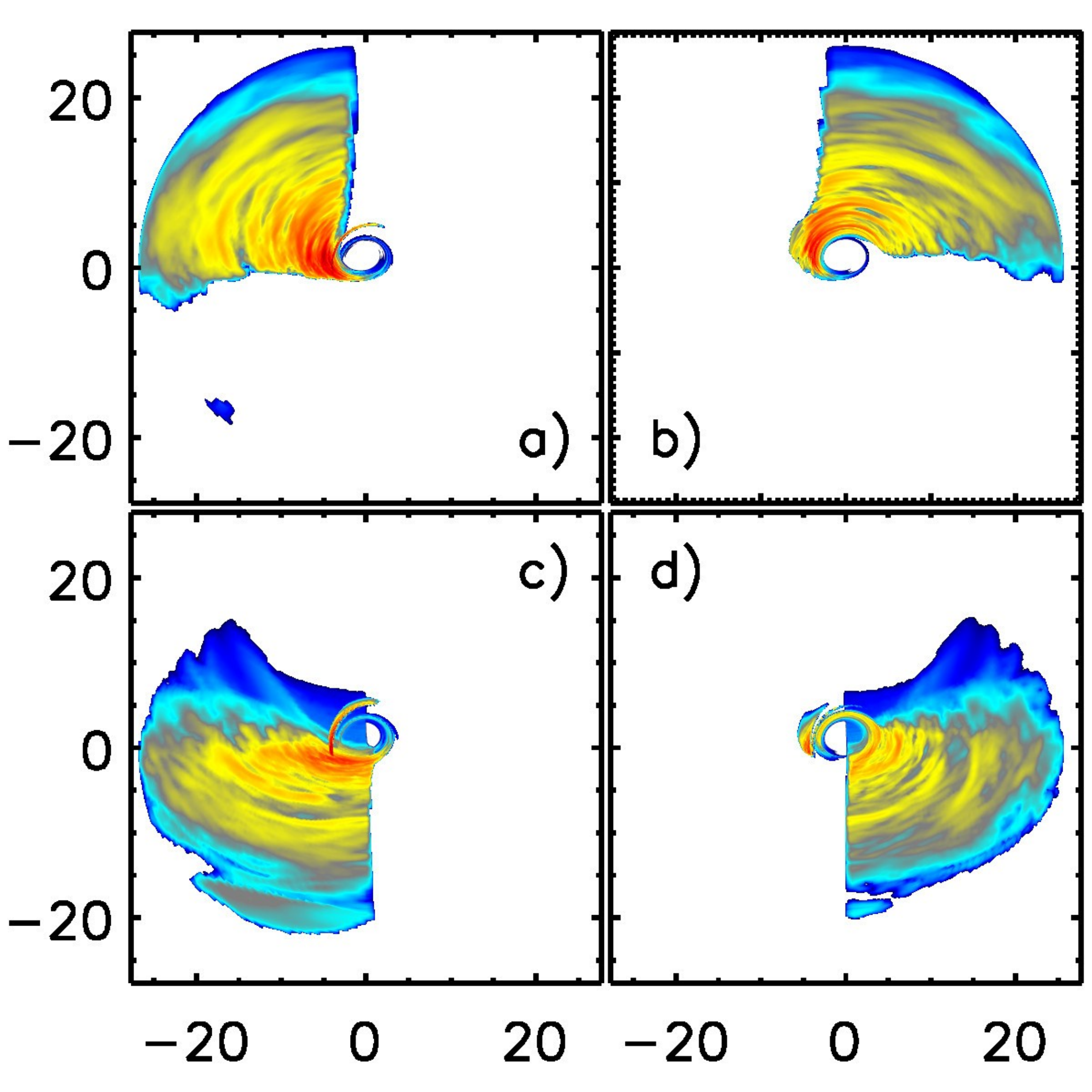
Degeneracy Explanation



Degeneracy Explanation

$\alpha_a > -2$

$\alpha_b > -2$



$\alpha_c < -2$

$\alpha_d < -2$

$\theta \sim 0^o$
$\alpha_i \simeq 2$

Summary & Conclusions

- Closer to ab initio calculations of accretion disk dynamics
- Magnetic stress is important within ISCO
- Stress does not vanish with disk height (at least for $a = 0$)
- Dissipation variability approximates observed coronal variability

- What about
  … other spins?
  … other cooling models?
    H = const. , H = H(t,r) Hysterisis? State Transitions?
  … other initial magnetic field topologies?
  … radiation pressure? (ugh)

- Near-merger BBH Disks …
  …. are magnetized and different from gap-forming hydro disks
    (magnetic stress can work over extended regions unlike visc.)
  … most likely will **not** have large gaps
  … will most likely be bright and variable before merger